

Properties of flat-spectrum radio-loud narrow-line Seyfert 1 galaxies[★]

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ABSTRACT

We have conducted a multiwavelength survey of 42 radio loud narrow-line Seyfert 1 galaxies (RLNLS1s), selected by searching among all the known sources of this type and omitting those with steep radio spectra. We analyse data from radio frequencies to X-rays, and supplement these with information available from online catalogues and the literature in order to cover the full electromagnetic spectrum. This is the largest known multiwavelength survey for this type of source. We detected 90% of the sources in X-rays and found 17% at γ rays. Extreme variability at high energies was also found, down to timescales as short as hours. In some sources, dramatic spectral and flux changes suggest interplay between a relativistic jet and the accretion disk. The estimated masses of the central black holes are in the range $\sim 10^{6-8} M_{\odot}$, lower than those of blazars, while the accretion luminosities span a range from ~ 0.01 to ~ 0.49 times the Eddington limit, with an outlier at 0.003, similar to those of quasars. The distribution of the calculated jet power spans a range from $\sim 10^{42.6}$ to $\sim 10^{45.6}$ erg s⁻¹, generally lower than quasars and BL Lac objects, but partially overlapping with the latter. Once normalised by the mass of the central black holes, the jet power of the three types of active galactic nuclei are consistent with each other, indicating that the jets are similar and the observational differences are due to scaling factors. Despite the observational differences, the central engine of RLNLS1s is apparently quite similar to that of blazars. The historical difficulties in finding radio-loud narrow-line Seyfert 1 galaxies might be due to their low power and to intermittent jet activity.

Key words. galaxies: Seyfert – galaxies: jets – quasars: general – BL Lacertae objects: general

1. Introduction

An important new discovery made with the Large Area Telescope (LAT, Atwood et al. 2009) on board the *Fermi* Gamma-ray Space Telescope is the high-energy gamma-ray emission from radio-loud narrow-line Seyfert 1 galaxies (RLNLS1s, Abdo et al. 2009a,b,c; Foschini et al. 2010). Narrow-line Seyfert 1 galaxies (NLS1s) are a well-known class of active galactic nuclei (AGNs), but they are usually considered to be radio-quiet (e.g. Ulvestad et al. 1995; Moran 2000; Boroson 2002). Thus, the first discoveries of RLNLS1s (e.g. Remillard et al. 1986; Grupe et al. 2000; Oshlack et al. 2001; Zhou et al. 2003) seemed to be exceptions, rather than the tip of an iceberg. The early surveys revealed only a handful of objects: 11 by Zhou & Wang (2002) and Komossa et al. (2006a), and 16 by Whalen et al. (2006). Williams et al. (2002) analysed 150 NLS1s from the *Sloan Digital Sky Survey* (SDSS) Early Data Release, and only a dozen (8%) were detected at radio frequencies and only two (1.3%) are radio loud, i.e. with the ratio between radio and optical flux densities greater than 10. One source is also in the present sample (J0948 + 0022, Zhou et al. 2003), while we have discarded the other (J1722 + 5654, Komossa et al. 2006b) because of its steep radio index (see Sect. 2). Most of the mildly radio-loud NLS1 galaxies of Komossa et al. (2006a) are steep-spectrum sources, and do not show indications of beaming, while three sources are more similar to blazars. In terms of

their optical emission-line properties and black hole masses, the RLNLS1s are similar to the radio-quiet NLS1 (RQNLS1) population as a whole. A larger study by Zhou et al. (2006) based on SDSS Data Release 3 resulted in a sample of 2011 NLS1s, about 14% of all the AGNs with broad emission lines. The fraction detected in the radio is 7.1%, similar to what was found by Williams et al. (2002). From this subsample, Yuan et al. (2008) culled 23 RLNLS1s with radio loudness greater than 100 and found that these sources are characterised by flat radio spectra. Detection of flux and spectral variability and their characteristic spectral energy distributions (SEDs) suggest a blazar-like nature.

In 2009, detection at high-energy γ rays by Abdo et al. (2009a,c) revealed beyond any reasonable doubt the existence of powerful relativistic jets in RLNLS1s and brought this poorly known class of AGNs into the spotlight (see Foschini 2012a for a recent review). An early survey including gamma-ray detections (after 30 months of *Fermi* operations) was carried out by Foschini (2011a). Forty-six RLNLS1s were found, of which seven were detected by *Fermi*. Of 30 RQNLS1 that served as a control sample, none were detected at γ rays. Additional multi-wavelength (MW) data, mostly from archives, were employed in this survey; specifically, X-ray data from ROSAT were used, but yielded a detection rate of only about 60%.

To improve our understanding of RLNLS1s, we decided to perform a more extended and detailed study. First, we have revised the sample selection (see Sect. 2), resulting in 42 RLNLS1s. We focus here on the population that is likely beamed (i.e. where the jet is viewed at small angles); a parallel

[★] Tables 4–9 and Figs. 8–13 are available in electronic form at <http://www.aanda.org>

study on the search for the parent population (i.e. with the jet viewed at large angles) is ongoing (Berton et al., in prep.). We therefore exclude from this study RLNLS1s with steep radio spectral indices, although we keep the sources with no radio spectral index information. We requested specific observations with *Swift* and *XMM-Newton* to improve the X-ray detection rate, which is now at 90%. Observations with these satellites were also accompanied by ultraviolet observations to study the accretion disk emission. Optical spectra were mostly taken from the SDSS archives and from the literature. For two sources, new optical spectra were obtained at the Asiago Astrophysical Observatory (Italy). New radio observations, particularly from monitoring campaigns on the γ -ray detected RLNLS1s, supplemented the archival data. More details on radio monitoring programs at Effelsberg/Pico Veleta and Metsähovi will be published separately (Angelakis et al. 2015; Lähteenmäki et al., in prep.). Some preliminary results from the present work have already been presented by Foschini et al. (2013).

To facilitate comparison with previous work, we adopt the usual Λ CDM cosmology with a Hubble–Lemaître constant $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_\Lambda = 0.73$ (Komatsu et al. 2011). We adopt the flux density and spectral index convention $S_\nu \propto \nu^{-\alpha_\nu}$.

2. Sample selection

The number of RLNLS1s known today is quite small compared to other classes of AGNs. We selected all the sources found in previous surveys (Zhou & Wang 2002; Komossa et al. 2006a; Whalen et al. 2006; Yuan et al. 2008) and from individual studies (Grupe et al. 2000; Oshlack et al. 2001; Zhou et al. 2003, 2005, 2007; Gallo et al. 2006) that meet the following criteria:

- Optical spectrum with an $H\beta$ line width $FWHM(H\beta) < 2000 \text{ km s}^{-1}$ (Goodrich 1989) with tolerance $+10\%$, a line-flux ratio $[O\text{III}]/H\beta < 3$, and clear broad Fe II emission blends (Osterbrock & Pogge 1985).
- Radio loudness $RL = S_{\text{radio}}/S_{\text{optical}} > 10$, where S_{radio} is the flux density at 5 GHz and S_{optical} is the optical flux density at 440 nm. In cases where 5 GHz fluxes are not available, we used other frequencies – generally 1.4 GHz – under the hypothesis of a flat radio spectrum, (i.e. $\alpha_r \approx 0$).
- Flat or inverted radio spectra ($\alpha_r < 0.5$, within the measurement errors), in order to select jets viewed at small angles. Sources with steep radio spectra (corresponding to jets viewed at large angles) are the subject of another survey (Berton et al., in prep.). Sources without spectral information and with only a radio detection at 1.4 GHz are included in our sample.

Radio loudness was recalculated on the basis of more recent data from Foschini (2011a), leading to some sources from Whalen et al. (2006) being reclassified as radio loud or radio quiet. Given the variability of the radio emission, we decided to keep all the sources which were classified as radio loud at least in one of the two samples. The resulting list of 42 sources studied in the present work is displayed in Table 1. For each source, we searched all the data available from radio to γ rays (see Sect. 3). It is worth noting that in this work we do not make a distinction between quasars and Seyfert galaxies, although most of the sources of the present samples are sufficiently luminous to be classified as quasars. We adopt the general acronym RLNLS1s for all the sources in the sample.

We also note that there has been some doubt about the classification of J2007–4434 as NLS1 because of its weak

Fe II emission: Komossa et al. (2006a) proposed a classification as narrow-line radio galaxy, while Gallo et al. (2006) argued that since there is no quantitative criterion on the intensity of Fe II, the source can be considered to be a genuine RLNLS1. We follow the latter interpretation and include J2007–4434 in our sample.

To facilitate comparison with blazars, we selected a sample of 57 flat-spectrum radio quasars (FSRQs) and 31 BL Lac objects, all detected by *Fermi*/LAT (Ghisellini et al. 2009, 2010; Tavecchio et al. 2010, and references therein). This sample was built by selecting all the sources in the LAT Bright AGN Sample (LBAS, Abdo et al. 2009d) with optical-to-X-ray coverage with *Swift* and information about masses of the central black holes and jet power. However, those works do not contain all the information we need to make a complete broad-band comparison with the present set of RLNLS1s. Therefore, we supplemented the published data in the cited works with information from online catalogues, specifically radio data at 15 GHz from the MOJAVE Project (Lister et al. 2009, 2013), ultraviolet fluxes from *Swift*/UVOT extracted from the Science Data Center of the Italian Space Agency (ASI-ASDC¹), and X-ray fluxes from the *Swift* X-ray Point Sources catalogue (1XSPS, Evans et al. 2014).

3. Data analysis and software

We retrieved all the publicly available observations done by *Swift* (Gehrels et al. 2004) and *XMM-Newton* (Jansen et al. 2001) on 2013 December 9. Data analysis was performed by following standard procedures as described in the documentation for each instrument.

In the case of *Swift* we used HEASoft v.6.15 with the calibration data base updated on 2013 Dec. 13. We analysed data of the X-Ray Telescope (XRT, Burrows et al. 2005) and the Ultraviolet Optical Telescope (UVOT, Roming et al. 2005). XRT spectral counts were rebinned to have at least 20–30 counts per bin in order to apply the χ^2 test. When this was not possible, we applied the unbinned likelihood (Cash 1979). We adopted power-law and broken power-law models. The need for the latter was evaluated by using the f -test (cf. Protassov et al. 2002) with a threshold $>99\%$. The observed magnitudes (Vega System) of UVOT were dereddened according to Cardelli et al. (1989) and converted into physical units by using zero points from *Swift* calibration data base. All the sources are point-like, and therefore we consider the emission from the host galaxy to be negligible; only J0324+3410 in the V filter displayed some hint of host galaxy, which was properly subtracted. We did not analyse the Burst Alert Telescope (BAT, Barthelmy et al. 2005) data because the average fluxes of RLNLS1s in hard X-rays are well below the instrument sensitivity. Indeed, by looking at the two available catalogues built on BAT data, we found only one detection of J0324+3410 in both the 70-month survey of the *Swift*/BAT team (Baumgartner et al. 2013) and the Palermo 54-month catalogue (Cusumano et al. 2010). J0324+3410 was first detected by Foschini et al. (2009) by integrating all the available direct observations performed during the period 2006–2008 (total exposure ~ 53 ks). There is also another detection of J0948+0022 in the Palermo catalogue, but not confirmed by Baumgartner et al. (2013). We did not include this information in the present work. *Swift* results are summarised in Tables 5 and 6.

In the case of *XMM-Newton*, we analysed data of the European Photon Imaging Camera (EPIC) pn (Strüder et al. 2001) and MOS (Turner et al. 2001) detectors. We adopted

¹ <http://www.asdc.asi.it/>

Table 1. Sample of RLNLS1s.

Name	Alias	α	δ	z	N_{H}	$FWHM \text{H}\beta$	$S_{1.4 \text{ GHz}}$
J0100 – 0200	FBQS J0100 – 0200	01:00:32.22	–02:00:46.3	0.227	4.12	920	6.4
J0134 – 4258	PMN J0134 – 4258	01:34:16.90	–42:58:27.0	0.237	1.69	930	55.0*
J0324 + 3410	1H 0323 + 342	03:24:41.16	+34:10:45.8	0.061	12.0	1600	614.3**
J0706 + 3901	FBQS J0706 + 3901	07:06:25.15	+39:01:51.6	0.086	8.27	664	5.6
J0713 + 3820	FBQS J0713 + 3820	07:13:40.29	+38:20:40.1	0.123	6.00	1487	10.4
J0744 + 5149	NVSS J074402 + 514917	07:44:02.24	+51:49:17.5	0.460	4.83	1989	11.9
J0804 + 3853	SDSS J080409.23 + 385348.8	08:04:09.24	+38:53:48.7	0.211	5.26	1356	2.9
J0814 + 5609	SDSS J081432.11 + 560956.6	08:14:32.13	+56:09:56.6	0.509	4.44	2164	69.2
J0849 + 5108	SDSS J084957.97 + 510829.0	08:49:57.99	+51:08:28.8	0.584	2.97	1811	344.1
J0902 + 0443	SDSS J090227.16 + 044309.5	09:02:27.15	+04:43:09.4	0.532	3.10	2089	156.6
J0937 + 3615	SDSS J093703.02 + 361537.1	09:37:03.01	+36:15:37.3	0.179	1.22	1048	3.6
J0945 + 1915	SDSS J094529.23 + 191548.8	09:45:29.21	+19:15:48.9	0.284	2.16	<2000	17.2
J0948 + 0022	SDSS J094857.31 + 002225.4	09:48:57.29	+00:22:25.6	0.585	5.55	1432	107.5
J0953 + 2836	SDSS J095317.09 + 283601.5	09:53:17.11	+28:36:01.6	0.658	1.25	2162	44.6
J1031 + 4234	SDSS J103123.73 + 423439.3	10:31:23.73	+42:34:39.4	0.376	1.01	1642	16.6
J1037 + 0036	SDSS J103727.45 + 003635.6	10:37:27.45	+00:36:35.8	0.595	5.07	1357	27.2
J1038 + 4227	SDSS J103859.58 + 422742.2	10:38:59.59	+42:27:42.0	0.220	1.50	1979	2.8
J1047 + 4725	SDSS J104732.68 + 472532.0	10:47:32.65	+47:25:32.2	0.798	1.31	2153	734.0
J1048 + 2222	SDSS J104816.58 + 222239.0	10:48:16.56	+22:22:40.1	0.330	1.51	1301	1.2
J1102 + 2239	SDSS J110223.38 + 223920.7	11:02:23.36	+22:39:20.7	0.453	1.22	1972	2.0
J1110 + 3653	SDSS J111005.03 + 365336.3	11:10:05.03	+36:53:36.1	0.630	1.85	1300	18.6
J1138 + 3653	SDSS J113824.54 + 365327.1	11:38:24.54	+36:53:27.0	0.356	1.82	1364	12.5
J1146 + 3236	SDSS J114654.28 + 323652.3	11:46:54.30	+32:36:52.2	0.465	1.42	2081	14.7
J1159 + 2838	SDSS J115917.32 + 283814.5	11:59:17.31	+28:38:14.8	0.210	1.70	1415	2.2
J1227 + 3214	SDSS J122749.14 + 321458.9	12:27:49.15	+32:14:59.0	0.137	1.37	951	6.5
J1238 + 3942	SDSS J123852.12 + 394227.8	12:38:52.15	+39:42:27.6	0.623	1.42	910	10.4
J1246 + 0238	SDSS J124634.65 + 023809.0	12:46:34.68	+02:38:09.0	0.363	2.02	1425	37.0
J1333 + 4141	SDSS J133345.47 + 414127.7	13:33:45.47	+41:41:28.2	0.225	0.74	1940	2.5
J1346 + 3121	SDSS J134634.97 + 312133.7	13:46:35.07	+31:21:33.9	0.246	1.22	1600	1.2
J1348 + 2622	SDSS J134834.28 + 262205.9	13:48:34.25	+26:22:05.9	0.918	1.17	1840	1.6
J1358 + 2658	SDSS J135845.38 + 265808.5	13:58:45.40	+26:58:08.3	0.331	1.56	1863	1.8
J1421 + 2824	SDSS J142114.05 + 282452.8	14:21:14.07	+28:24:52.2	0.538	1.28	1838	46.8
J1505 + 0326	SDSS J150506.47 + 032630.8	15:05:06.47	+03:26:30.8	0.409	4.01	1082	365.4
J1548 + 3511	SDSS J154817.92 + 351128.0	15:48:17.92	+35:11:28.4	0.479	2.37	2035	140.9
J1612 + 4219	SDSS J161259.83 + 421940.3	16:12:59.83	+42:19:40.0	0.234	1.29	819	3.4
J1629 + 4007	SDSS J162901.30 + 400759.9	16:29:01.31	+40:07:59.6	0.272	1.06	1458	12.0
J1633 + 4718	SDSS J163323.58 + 471858.9	16:33:23.58	+47:18:59.0	0.116	1.77	909	62.6
J1634 + 4809	SDSS J163401.94 + 480940.2	16:34:01.94	+48:09:40.1	0.495	1.66	1609	7.5
J1644 + 2619	SDSS J164442.53 + 261913.2	16:44:42.54	+26:19:13.2	0.145	5.12	1507	87.5
J1709 + 2348	SDSS J170907.80 + 234837.6	17:09:07.82	+23:48:38.2	0.254	4.12	1827	1.6
J2007 – 4434	PKS 2004 – 447	20:07:55.18	–44:34:44.3	0.240	2.93	1447	791.0***
J2021 – 2235	IRAS 20181 – 2244	20:21:04.38	–22:35:18.3	0.185	5.54	460	24.9**

Notes. Columns: (1) name of the source as used in the present work; (2) other name often found in the literature; (3) right ascension (J2000); (4) declination (J2000); (5) redshift from SDSS or NED; (6) galactic absorption column density [10^{20} cm^{-2}] from Kalberla et al. (2005); (7) full-Width Half Maximum of broad $\text{H}\beta$ emission line [km s^{-1}]; (8) peak radio flux density at 1.4 GHz from VLA/FIRST (Becker et al. 1995) or from the nearest frequency available [mJy]. The coordinates were mostly from the VLA/FIRST survey; when missing, we referred to NED. (*) 4.85 GHz; Grupe et al. (2000); (**) VLA/NVSS, Condon et al. (1998); (***) 1.4 GHz, ATCA, Gallo et al. (2006).

the Science Analysis Software v.13.5.0 with the calibration data base updated on 2013 December 19. We excluded time periods with high-background by following the prescriptions of Guainazzi et al. (2013). The spectral modelling was done as for *Swift*/XRT. *XMM-Newton* results are summarised in Table 5.

3.1. Optical data

Optical spectra were retrieved for 32/42 sources from SDSS DR9 database (Ahn et al. 2012), downloaded from NED (3/42), or extracted from figures published in the literature (2/43). Two sources, J0324+3410 and J0945+1915, were observed with the 1.22 m telescope of the Asiago Astrophysical Observatory between 2013 December and 2014 January, using

the Boller & Chivens spectrograph with a 300 mm^{-1} grating. The instrumental resolution was $R \approx 700$, and the spectra covered the wavelength range between 3200 and 8000 Å with a dispersion of $2.3 \text{ \AA pixel}^{-1}$. The slit was oriented at $\text{PA} = 90^\circ$, with an aperture of 4.25 arcsec, corresponding to 4.7 kpc for J0324+3410 and to 17.6 kpc for J0945+1915. The exposure time was $3 \times 1200 \text{ s}$ for the former and $9 \times 1200 \text{ s}$ for the latter. Data reduction was performed using the standard IRAF v.2.14.1 tasks: the overscan was subtracted instead of the bias in the pre-reduction steps and NeHgAr lamps were used for the wavelength calibration. Finally the extracted spectra were combined together (see Fig. 1).

We were unable to find any optical spectral data for three of the sources in our sample.

The optical spectra were corrected for redshift and Galactic absorption and a continuum fit was subtracted. The contribution

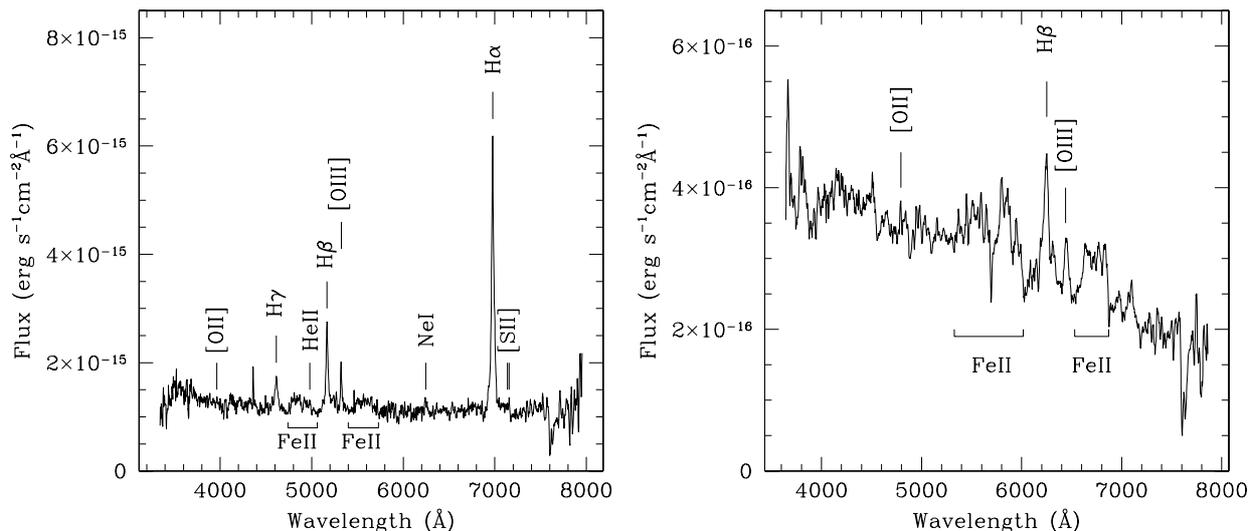


Fig. 1. Optical spectra of J0324+3410 (left panel) and J0945+1915 (right panel) taken from the Asiago Astrophysical Observatory 1.22 m telescope.

of the host galaxy in objects at $z > 0.1$ is typically less than 10% (Letawe et al. 2007). Given that the flux calibration uncertainty is typically around 20%, we assume that the host galaxy contribution is negligible. Indeed the spectra, as expected, do not show any sign of stellar absorption features, and the continuum appears to be dominated by the AGN. For objects at $z < 0.1$ (J0324+3410 and J0706+3901), we subtracted a template of a spiral galaxy bulge (Kinney et al. 1996) as a test, even if no stellar features were visible. Since we did not observe any significant change in the shape of $H\beta$, we proceeded without any host-galaxy subtraction. We focused on the $H\beta$ region between 4000 and 5500 Å. To subtract Fe II multiplets, we used a template properly created by using the online software² developed by Kovačević et al. (2010) and Shapovalova et al. (2012).

After Fe II subtraction, we decompose the $H\beta$ line into narrow and broad components, using the `ngaussfit` task of IRAF. We used three Gaussians to fit the profile, one to reproduce the narrow component, and two others for the broad component. Following Veron et al. (2001), we fixed the flux of the narrow component to be 1/10 of the [O III] $\lambda 5007$ line with the same velocity width. However, given that the gas which produces the [O III] line is often turbulent, its width can lead to an overestimate of the $H\beta$ narrow component. For this reason, when [O II] $\lambda 3727$ was clearly visible and much narrower than [O III] lines, we used its FWHM to fix the width of $H\beta$ (Greene & Ho 2005; Ho et al. 2009). In some case, the low signal-to-noise ratio (S/N) required a fit with just two Gaussians, one narrow and one broad. When necessary we also set the height of the narrow component as a free parameter. The line centre was always left as a free parameter.

In the case of J1348+2622, we used the Mg II $\lambda 2798$ for the black hole mass estimate as the $H\beta$ line it falls outside of the spectral range. As shown by Shen et al. (2008), mass estimates from these two lines are generally consistent.

Finally, we subtracted the narrow component and measuring both FWHM and the line dispersion σ only for the broad component. The results are presented in Table 2.

3.2. Radio data

Some of these sources were observed for other programs. 37 GHz data are from the 13.7 m telescope at Metsähovi (Finland), and MW observations were done at 100-m single dish telescope at Effelsberg (Germany, 2.64–42 GHz) and 30 m telescope at Pico Veleta (Spain, 86–142 GHz). More details about the Metsähovi, and Effelsberg/Pico Veleta observations on RLNLS1s will be published by Lähteenmäki et al. (in prep.) and Angelakis et al. (2015), respectively. Some of the data have already been published by Abdo et al. (2009a,b), Foschini et al. (2011a, 2012), Fuhrmann et al. (2011), and Angelakis et al. (2012a,b).

We also searched for publicly available observations in the VLBI calibrated data archives. 15 GHz data are from the MOJAVE database (Lister et al. 2009, 2013)³. VLBI results at frequencies below 15 GHz come from the VLBA and global VLBI astrometric and geodetic experiments (Beasley et al. 2002; Fomalont et al. 2003; Petrov et al. 2005, 2006, 2008; Kovalev et al. 2007; Piner et al. 2012; Pushkarev & Kovalev 2012). Calibrated visibility and image fits files are provided by the authors in the public database⁴.

We performed a standard CLEANing (Högbom 1974) and followed the model-fitting of the calibrated VLBI visibility data in `DiFmap` (Shepherd 1997). We preferred to use circular Gaussian components unless the use of elliptical components gave a better fit to the data. To ensure the quality of the fit, we compared Gaussian model parameters with the results of CLEAN. The total flux density and residual RMS appeared to be almost identical for the two cases. All of these sources have simple radio structure, so they are well-modelled by Gaussian components. The results are presented in Table 7.

3.3. Online catalogues and literature

We supplemented these data with information from online catalogues and literature. For γ rays, we mainly referred to Foschini (2011a), who reported the detection of 7 RLNLS1s with *Fermi*/LAT after 30 months of operations. When available,

² <http://servo.aob.rs/~jelena/>

³ <http://www.physics.purdue.edu/MOJAVE/>

⁴ <http://astrogeo.org>

Table 2. Mass and accretion luminosity estimated from optical data.

Source	σ_{line}	<i>FWHM</i>	$L_{\text{H}\beta}$	<i>M</i>	L_{disk}	$L_{\text{disk}}/L_{\text{Edd}}$	Method
J0100 – 0200	982	920	–	4.0	25.2	0.49	M
J0134 – 4258	1632	1241	2.77	7.1	8.61	0.09	L (Grupe et al. 2000)
J0324 + 3410	1791	1868	0.53	3.6	1.50	0.03	A
J0706 + 3901	1839	1402	0.16	2.0	0.43	0.02	N
J0713 + 3820	2041	1901	9.28	21.2	31.0	0.11	N
J0744 + 5149	2122	1989	–	26.5	49.2	0.14	M
J0804 + 3853	1588	1523	2.74	6.7	8.51	0.10	S
J0814 + 5609	2759	2777	6.10	31.0	19.9	0.05	S
J0849 + 5108	1330	2490	1.32	3.2	3.94	0.09	S
J0902 + 0443	1491	1781	2.02	5.0	6.17	0.09	S
J0937 + 3615	1343	2192	0.59	2.1	1.68	0.06	S
J0945 + 1915	1730	2818	2.52	7.6	7.80	0.08	A
J0948 + 0022	1548	1639	3.73	7.5	11.8	0.12	S
J0953 + 2836	2407	2749	3.14	16.6	9.84	0.05	S
J1031 + 4234	3153	1822	2.08	22.9	6.37	0.02	S
J1037 + 0036	985	1776	1.75	2.0	5.31	0.20	S
J1038 + 4227	1615	1917	2.62	6.8	8.12	0.09	S
J1047 + 4725	1474	2237	5.80	8.6	18.9	0.17	S
J1048 + 2222	1742	718	1.39	5.7	4.16	0.06	S
J1102 + 2239	1940	2181	3.56	11.5	11.2	0.08	S
J1110 + 3653	1230	2081	1.07	2.4	3.14	0.10	S
J1138 + 3653	1231	1542	1.29	2.7	3.85	0.11	S
J1146 + 3236	1737	1977	3.79	9.5	12.0	0.10	S
J1159 + 2838	1907	2728	0.052	1.2	0.13	0.01	N
J1227 + 3214	694	1567	0.51	0.52	1.42	0.21	S
J1238 + 3942	940	1229	1.16	1.5	3.42	0.18	S
J1246 + 0238	1667	1756	1.80	5.9	5.45	0.07	S
J1333 + 4141	1589	2942	1.73	5.3	5.25	0.08	S
J1346 + 3121	1074	1503	0.64	1.4	1.83	0.10	S
J1348 + 2622	2192	3361	4.09	5.3	13.0	0.19	S (based on Mg II λ 2798)
J1358 + 2658	1471	1805	3.20	6.3	10.1	0.12	S
J1421 + 2824	1589	1724	7.18	11.2	23.7	0.16	S
J1505 + 0326	1409	1337	0.41	1.9	1.12	0.05	S
J1548 + 3511	1557	2217	4.37	8.3	14.0	0.13	S
J1612 + 4219	777	1200	0.87	0.88	2.53	0.22	S
J1629 + 4007	1246	1410	2.00	3.5	6.10	0.13	S
J1633 + 4718	945	931	0.36	0.79	0.98	0.10	S
J1634 + 4809	1856	1763	1.95	7.7	5.94	0.06	S
J1644 + 2619	1129	1486	0.51	1.4	1.42	0.08	S
J1709 + 2348	2377	1256	0.95	2.4	2.79	0.09	S
J2007 – 4434	1869	2844	0.0081	0.43	0.018	0.003	L (Oshlack et al. 2001)
J2021 – 2235	491	460	–	3.75	2.91	0.60	M

Notes. Columns: (1) name of the source; (2) line dispersion σ_{line} of the broad component of H β [km s $^{-1}$]; (3) FWHM of the broad component of H β [km s $^{-1}$]; (4) H β luminosity [10 42 erg s $^{-1}$]; (5) black hole mass [10 $^7 M_{\odot}$]; (6) disk luminosity [10 44 erg s $^{-1}$]; (7) disk luminosity [Eddington units]; (8) method adopted: A, from Asiago spectra; L, spectra from literature; M, derived from optical magnitudes; N, spectra downloaded from NED; S, from SDSS spectra.

we reported more recent published analyses (Foschini et al. 2012; D’Ammando et al. 2013a,d; Paliya et al. 2014). No new detections have been claimed to date after Foschini (2011a). Therefore, for the non-detected sources in the present sample, we indicated the upper limit of $\sim 10^{-9}$ ph cm $^{-2}$ s $^{-1}$ as from the *Fermi*/LAT performance web page⁵, which is the minimum detectable ($TS = 25$) flux above 100 MeV over a period of 4 years for a source with a power-law shaped spectrum with a spectral index $\alpha = 1$. A summary of γ ray characteristics found in literature is shown in Table 4.

For X-rays, we searched for missing sources in the *Chandra* X-Assist (CXA, Ptak & Griffiths 2003) catalogue v.4 and *XMM-Newton* Slew Survey Clean Sample v.1.5 (XSS,

Saxton et al. 2008). The two catalogues provide X-ray detections in different energy bands: 0.5–8 keV for the former, and 0.2–12 keV for the latter. The fluxes were then converted into the 0.3–10 keV band by using WebPIMMS⁶ and a fixed photon index value $\Gamma = 2$ ($\alpha = 1$). Some sources were not observed by any of the above-cited satellites. In those cases, we calculated an upper limit by using the detection limit of the ROSAT All-Sky Survey (RASS, Voges et al. 1999, 2000).

At infrared/optical/ultraviolet wavelengths, in addition to the *Swift*/UVOT data presented here, we used SDSS-III data release 9 (Ahan et al. 2012) and 2MASS (Cutri et al. 2003). Only one source, J2021–2235, remained without optical coverage from either *Swift*/UVOT or SDSS, but we found *B* and

⁵ http://www.slac.stanford.edu/exp/glast/groups/canda/lat_Performance.htm

⁶ <http://heasarc.gsfc.nasa.gov/Tools/w3pimms.html>

R magnitudes in the US Naval Observatory B1 catalogue (Monet et al. 2003).

We also searched the WISE all-sky catalogue (Wright et al. 2010) for photometric data at mid-IR wavelengths (between 3.4 and 22 μm). In particular, we have used the last version of the catalogue, the *AllWISE* data release (November 2013). All the RLNLS1s of the sample are detected ($S/N > 3$) in the WISE survey at 3.4 and 4.6 μm (*W1* and *W2* bands, respectively) while 41 and 37 objects are detected also at 12 μm (*W3* band) and 22 μm (*W4* band) respectively. The observed magnitudes have been converted into monochromatic flux densities assuming a power-law spectrum with $\alpha = 2$. For the sources not detected or detected with a $S/N < 3$, we have calculated the 3σ upper limit on the flux density.

At radio frequencies, in addition to the above cited programs (see Sect. 3.2), we have taken all available data from the NED⁷ and HEASARC⁸ archives.

4. Observational characteristics

4.1. Gamma rays

We found in the available literature 7/42 detections at high-energy γ rays (17%) sources. Specifically, they are:

- J0948+0022, the first RL-NLS1 to be detected in γ rays (Abdo et al. 2009a,b; Foschini et al. 2010).
- J0324+3410, J1505+0326, and J2007–4434, which were detected after the first year of *Fermi* operations (Abdo et al. 2009c).
- J0849+5108, which was detected because of an outburst in 2010 (Foschini 2011a; D’Ammando et al. 2012).
- J1102+2239, J1246+0238 (Foschini 2011a).

The spectral indices are generally steep, with a weighted average of $\alpha_\gamma = 1.6 \pm 0.3$ (median 1.7), but there is one interesting case with harder spectrum: J0849+5108 with $\alpha_\gamma = 1.0\text{--}1.18$ (Tables 4 and 8). The average values for blazars as measured by *Fermi*/LAT (Ackermann et al. 2011) are 1.4 ± 0.2 for FSRQs, $\alpha = 1.2 \pm 0.1$ for low-synchrotron peak BL Lacs, and $\alpha = 0.9 \pm 0.2$ for high-synchrotron peak BL Lacs. Therefore, we conclude that the spectral characteristics of RLNLS1s are generally similar to those of FSRQs.

Short timescale variability for factor-of-two flux changes is also reported by some authors. Specifically, Foschini (2011a) reported intraday variability for J0948+0022 and J1505+0326, while Paliya et al. (2014) found 3 h variability of J0324+3410 during its outburst of 2013 August 28 to 2013 September 1 (see Table 9).

4.2. X-rays

About 90% of the sources in the present sample (38/42) are detected in X-rays (see Table 5). The average spectral index in the 0.3–10 keV energy range is $\alpha_X = 1.0 \pm 0.5$, with a median value of 0.8 (see Table 8), as compared with the values of 0.58 (FSRQs), 1.3 (BL Lac objects), 1.1 (BLS1s), and 1.7 (RQNLS1s). These values were calculated from the samples of γ -ray blazars from Ghisellini et al. (2009, 2010) and Tavecchio et al. (2010) and radio-quiet Seyferts from Grupe et al. (2010). The corresponding distributions are displayed in

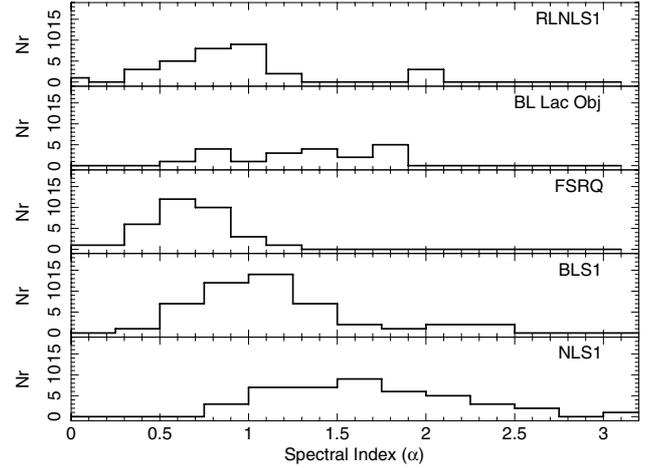


Fig. 2. X-ray (0.3–10 keV) spectral index distributions for the present sample of RLNLS1s; BL Lac objects and FSRQs are from Ghisellini et al. (2009, 2010) and Tavecchio et al. (2010); BLS1s and RQNLS1s are from Grupe et al. (2010).

Fig. 2. The average spectral indices for the individual sources (see Table 8) are $\alpha < 1$ in 23/42 cases and $\alpha \geq 1$ in 12/42 cases. In 7/42 cases, the spectral index is near the boundary.

We note that the X-ray spectral indices of RLNLS1s are similar to those of BLS1s, and usually harder than those of RQNLS1s. However, when compared to blazars, RLNLS1s are between the average values of FSRQs and BL Lac objects. From inspection of the SEDs (see Sect. 7), it seems that the X-ray emission of RLNLS1s could be due either to inverse-Compton (IC) radiation from a relativistic jet or from the corona of the accretion disk. This could explain why the average spectral index is softer than that of FSRQs, where the X-ray emission is dominated by the IC from the jet (see, for example, Ghisellini et al. 2010).

In the case of four sources, there were multiple observations with sufficient exposure for individual detections (Table 5). We therefore searched for any correlation between flux and spectral slope. No significant trend was found. It is interesting to compare with radio-quiet Seyferts, where a correlation between 2–10 keV flux and the spectral index was found, indicating a steepening of the spectral shape as the flux increases (Markowitz et al. 2003). Some interesting episodes were described in the case of J0324+3410 by Foschini et al. (2009), Foschini (2013), and Tibolla et al. (2013): the source has generally a soft spectral index, typical of NLS1s, but sometimes – as the jet became active – the X-ray spectrum displays a break at a 2–3 keV and a hard tail appears (see also Paliya et al. 2014). Similar behaviour has been observed in another RLNLS1, PKS 0558–504⁹, where ASCA observed a hardening of the spectral index during an outburst, changing from ~ 1.2 to ~ 0.9 (Wang et al. 2001).

In the case of J0324+3410, with more data (Table 5), there is no evident trend to link the change in flux with a change of the spectral slope. Although the epochs with a hard tail are concentrated in a high-flux region, there are also observations with similar fluxes that can be fit satisfactorily by a single power-law model. It is worth noting that there might be an instrumental bias: *Swift*/XRT has a small effective area at energies ≥ 7 keV (Romano et al. 2005) and, therefore, the detection of the hard

⁹ This source is not in the present sample because it has a steep radio spectral index. It is included in the sample studied by Berton et al. (in prep.).

⁷ <http://ned.ipac.caltech.edu/>

⁸ <http://heasarc.gsfc.nasa.gov/>

tail could depend on the exposure time at similar flux levels. Indeed, the exposures in the present data set ranged from 1.3 to 8.8 ks (see Table 5) and the spectral shape at shorter exposures – having low statistics at high energies – could be fit with just a power-law model with an index harder than usual. An observing campaign with a satellite like *XMM-Newton* for example, carrying X-ray instruments with a large effective area above 7 keV, could effectively monitor the spectral changes (see below the example of J0948+0022).

Many sources of the present sample were included in previous surveys by Komossa et al. (2006a) and Yuan et al. (2008). In these studies, the X-ray characteristics were measured from ROSAT data. Komossa et al. (2006a) found spectral indices in the range 0.9–3.3, while Yuan et al. (2008) measured values between 0.37 and 2.36. Particularly, the spectral index of J0948+0022 was measured as 1.6 ± 1.8 by Zhou et al. (2003), ~ 1.2 by Komossa et al. (2006a), and 1.26 ± 0.64 by Yuan et al. (2008). The source is also present in the Williams et al. (2002) sample, who reported $\alpha_X = 1.8 \pm 0.5$, again on the basis of ROSAT observations. In our case, both *Swift* and *XMM-Newton* indicate a harder spectral index (average $\alpha_X \sim 0.56$) that remains unchanged with flux variations (Table 5). This is in agreement with previous studies and MW campaigns (Abdo et al. 2009a,b; Foschini et al. 2011a, 2012). A study based on a long one-orbit *XMM-Newton* observation revealed the presence of a soft X-ray excess (D’Ammando et al. 2014; Bhattacharyya et al. 2014), which is confirmed in the present study. The break energy is between $1.72^{+0.09}_{-0.11}$ keV (D’Ammando et al. 2014) and ~ 1.2 keV (Bhattacharyya et al. 2014; see also our analysis in Table 5). The low-energy spectral index is between 1.1 and 1.3, while the high-energy power-law has a slope 0.5–0.6. There is also a *Swift* observation the same day of that of *XMM-Newton* (2011 May 28) and we tried also to fit these data with a broken power-law model. We found $\alpha_1 = \Gamma_1 - 1 = 1.8^{+1.4}_{-0.8}$, $\alpha_2 = \Gamma_2 - 1 = 0.5 \pm 0.2$, and $E_{\text{break}} = 0.9 \pm 0.3$ keV ($\chi^2 = 5.4$ for 7 dof, not reported in Table 5). However, according to our threshold defined in Sect. 3, the broken power-law model is not statistically preferred over the single power-law model (95% vs. a threshold of 99%). Therefore, we conclude that the presence or absence of a soft X-ray excess is related more to an instrumental bias rather than to an effective change of the AGN. ROSAT, having a bandpass of 0.1–2.4 keV, is biased toward soft X-ray sources and therefore only captures the soft excess. *Swift*, with a wider energy band (0.3–10 keV) and snapshot observations, measured an average of both the soft excess and the hard tail. *XMM-Newton*, still operating in the 0.3–10 keV band, detected both the soft excess and the hard tail because of the longer exposure and larger effective area. Both D’Ammando et al. (2014) and Bhattacharyya et al. (2014) concluded that the excess at low energies could be due to the accretion disk/corona system, as is the usual case for many RQNLs1s (e.g. Leighly 1999; Foschini et al. 2004; Grupe et al. 2010).

We note that also the blazar 3C 273 displays such a soft excess and there is a correlation between the low activity of the relativistic jet and the emergence of the thermal component in the soft X-rays, which was interpreted as a signature of the jet-disk connection (Grandi & Palumbo 2004; Foschini et al. 2006). In the present case, the instrumental bias prevents any conclusion about the X-ray component, but the optical component offers some hints that support the above hypothesis (see Sect. 7).

The case of J2007–4434 is different. The X-ray spectrum as observed by *XMM-Newton* on 2004 April 11 shows a soft excess and a hard tail. Gallo et al. (2006) favoured the hypothesis of an accretion disk corona to generate the excess low-energy flux,

while Foschini et al. (2009), on the basis of different variability characteristics (16% and <8% in the 0.2–1 keV and 2–10 keV energy bands, respectively), suggested a similarity with low-energy peaked BL Lac objects (i.e. the low-energy component is the tail of the synchrotron emission). This seems to be confirmed by the analysis and modelling of the SED (Abdo et al. 2009c; Paliya et al. 2013b; see also the Sect. 7). In the present work, we find two more *XMM-Newton* archival observations performed in 2012 (May 1 and October 18): in both cases, there was no low-energy excess and the spectra were fit with a single power-law model with spectral index $\alpha \approx 0.7$. It is worth noting that the flux was about one third that of the 2004 observation, when the soft excess was detected.

In all the other cases, ROSAT observations reported by Komossa et al. (2006a) and Yuan et al. (2008) are generally confirmed.

The search for variability on short timescales resulted in many significant detections of intraday variability, with flux changes greater than 3σ (Table 9). There are hour timescales for J0134–4258, J0324+3410, J0948+0022, J1629+4007, and J2007–4434. It is worth noting that the measurements reported in Table 9 were only made from *Swift*/XRT data. We also analysed *XMM-Newton* data and find variability on minute timescales (down to ~ 2 min with flux change at the 11σ level in the case of J0948+0022). However, we note that all *XMM-Newton* observations are affected by soft-proton flares: although we corrected for both the anomalous particle and photon backgrounds, we noted that minute-scale variability is detected near periods of the light curve that are excised because of high particle background. In addition, there is no confirmation of such short timescale variations in the *Swift* data, but it is worth noting that Itoh et al. (2013) found similar values from optical observations. These findings therefore require a much more careful dedicated analysis and confirmation with other instruments less affected by grazing-incidence particle background (i.e. from X-ray satellites in low-Earth orbit).

The hour timescales found are much shorter than those expected in case of changing obscuration, which could be ~ 10 h in the most extreme case of NGC 1365 (see the review by Bianchi et al. 2012). In addition, fits of X-ray spectra do not require iron lines or obscuration in addition to the Galactic column, as expected from radio-loud AGNs, in contrast to radio-quiet AGNs (e.g. Reeves et al. 1997). The exception seems to be J0324+341, where Abdo et al. (2009c) reported an unresolved iron line at $E_{\text{Fe}} = 6.5 \pm 0.3$ keV with equivalent width of 147 eV (see also Paliya et al. 2014). By integrating all the available *Swift* snapshots (with a total exposure time of 2.1×10^5 s), we basically confirm the previous measurements: $E_{\text{Fe}} = 6.5 \pm 0.1$ keV, equivalent width ~ 91 eV, and $\Delta\chi^2 = 13.1$ for two additional parameters (E_{Fe} and normalisation; we fixed σ_{Fe} to 0.1 keV). *XMM-Newton* observations of J0948+0022 and J1348+2622 only show a hint of an excess above 5 keV.

4.3. Ultraviolet, optical, and infrared frequencies

Spectral indices for ultraviolet and infrared frequencies were calculated by means of the two-point spectral index formula

$$\alpha_{12} = -\frac{\log(S_1/S_2)}{\log(\nu_1/\nu_2)}, \quad (1)$$

where S_1 and S_2 are the observed flux densities at frequencies ν_1 and ν_2 , respectively. In the ultraviolet, we use the *Swift*/UVOT observations, where $\nu_1 = 1.16 \times 10^{15}$ Hz and $\nu_2 = 1.47 \times 10^{15}$ Hz

refer to the *uvw1* ($\lambda = 2591 \text{ \AA}$) and *uvw2* ($\lambda = 2033 \text{ \AA}$) bands, respectively. For infrared frequencies, we adopt the extreme filters of WISE: $\nu_1 = 1.36 \times 10^{13} \text{ Hz}$ and $\nu_2 = 8.82 \times 10^{13} \text{ Hz}$, corresponding to *W4* and *W1* filters, respectively. When one of the two filters only has an upper limit, we referred to the closest other filter with a detection, either *W2* ($\nu = 2.50 \times 10^{13} \text{ Hz}$) or *W3* ($6.5 \times 10^{13} \text{ Hz}$). Optical spectral indices were measured by fitting the spectra over the range $\sim 3000\text{--}8000 \text{ \AA}$.

The average ultraviolet spectral index is $\alpha_{\text{uv}} = 0.7 \pm 1.4$ (median 0.7), and the values for the individual sources were measured from the integrated flux densities (Table 6). This spectral index can be compared with the values of 0.79 (median 0.61) for BLS1, and 0.85 (median 0.65) for NLS1 in the Grupe et al. (2010) sample, also based on *Swift*/UVOT observations. Ganguly et al. (2007) observed 14 radio-quiet low redshift ($z < 0.8$) quasars with *Hubble* Space Telescope (HST) in the range $1570\text{--}3180 \text{ \AA}$ and measured an average index of 0.87. Pian et al. (2005), also with HST over the range $\sim 1570\text{--}4780 \text{ \AA}$, observed 16 blazars and found $\alpha_{\text{uv}} \approx 1.16$. In a previous study on a larger sample of 47 radio-loud AGNs observed in the range $1200\text{--}3000 \text{ \AA}$ with International Ultraviolet Explorer (IUE), Pian & Treves (1993) found an average $\alpha_{\text{uv}} \approx 1$, with strong emission line quasars having $\alpha_{\text{uv}} \approx 1.38$, BL Lac objects with $\alpha_{\text{uv}} \approx 0.97$, and radio-weak BL Lacs with $\alpha_{\text{uv}} \approx 0.66$. For a control sample of 37 objects from the Palomar-Green (PG) bright quasar survey, an average $\alpha_{\text{uv}} \approx 0.84$ was found. At the level of average values, the present sample is in agreement with the values for radio-weak blazars, PG quasars, and radio-quiet Seyferts.

The average optical spectral index of the present sample of RLNLS1s is $\alpha_o = -1.0 \pm 0.8$ (median -0.8), in agreement with the previous surveys of RQ-NLS1 (Constantin & Shields 2003) and RLNLS1s (Komossa et al. 2006a; Yuan et al. 2008). A comparison with the optical slopes measured by Whalen et al. (2006) reveals similar slopes (particularly, Fig. 4 in Whalen et al. 2006), with some exceptions. For J0100–0200, Whalen et al. (2006) find $\alpha_o \approx -0.25$, while in the present work, we find (from SDSS photometry) $\alpha_o \approx 1.18$. Other cases of changes in the sign of the slope were J1159+2838 (from $\alpha_o \approx 0.19$ to -0.04) and J1421+2824 (from $\alpha_o \approx 0.18$ to -0.19). Vanden Berk et al. (2001) integrated the SDSS spectra of more than 2200 quasars and found an average $\alpha_o \approx -0.44$. They also note that by using only the low-redshift sources, the optical spectral index becomes steeper ($\alpha_o \approx -0.65$). Our average value ($\alpha_o \approx -1.03$) seems to be in agreement with this trend.

The average infrared spectral index as measured from WISE is $\alpha_{\text{IR}} = 1.3 \pm 0.3$ (median 1.3), as expected in the case of synchrotron emission from a relativistic jet (see Massaro et al. 2011; D’Abrusco et al. 2012; Raiteri et al. 2014). Figure 3 shows the distribution of WISE colours of the present sample compared with the blazar strip by Massaro et al. (2011) and the X-ray selected Type 1 and 2 AGNs by Mateos et al. (2012, 2013). While most of the RLNLS1s are in the blazar/AGN region, there are also some cases in the starburst region (cf. Fig. 1 in Massaro et al. 2011), suggesting the presence of intense star-formation activity (typical of NLS1s, Sani et al. 2010). Caccianiga et al. (2014) studied a steep-spectrum RLNLS1 (SDSS J143244.91+301435.3, which is not included in the present sample because of its steep radio spectrum), and found significant star-forming activity. In that case, since the jet is likely to be viewed at a large angle, it does not overwhelm the emission from the nearby environment. The RLNLS1s of the present sample were instead selected by their flat radio spectra, to extract the beamed population, and are hence dominated by

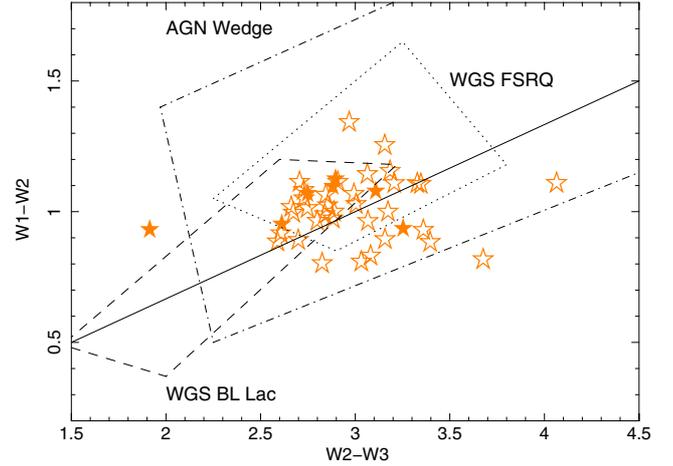


Fig. 3. WISE colours of the present sample of RLNLS1s (filled orange stars indicated the γ -ray detected RLNLS1s). Different characteristic regions are also plotted: the blazar WISE Gamma-ray Strip (WGS) for BL Lacs (dashed line) and FSRQs (dotted line), as defined by Massaro et al. (2012) and the AGN wedge (dot-dashed line) as defined by Mateos et al. (2012, 2013) for X-ray selected AGNs. The continuous line corresponded to a power-law emission ($S_\nu \propto \nu^{-\alpha}$) with α ranging from ~ -0.5 to $\sim +2.5$.

synchrotron emission. However, one source (J1505+0326) in the starburst region was detected in γ rays, and two γ -ray RLNLS1s have $W2 - W3 > 3$ (in the addition to the one cited earlier, there is also J0948+0022, still around the synchrotron line). Specifically, J0948+0022 and J1505+0326 were among the most γ -ray active RLNLS1s of the present sample, therefore this result could be counterintuitive (except for J0948+0022, which is still near the synchrotron line). The explanation is in the epochs of the WISE observations, performed between 2010 Jan. 7 (MJD 55 203) and Aug. 6 (MJD 55 414). Comparing with the γ -ray light curves displayed in Foschini et al. (2012), it is seen that J0948+0022 was almost undetected during the WISE observations. In the case of J1505+0326, D’Ammando et al. (2013a) detected the source by integrating over three-month time bins, but the flux in the first half of 2010 remained at low level of order $10^{-8} \text{ ph cm}^{-2} \text{ s}^{-1}$. Therefore, it is likely that these sources could have strong star-forming activity that is sometimes overwhelmed by synchrotron emission from the jet. We also note that another RLNLS1 of the present sample, J2021–2235, is classified as an ultraluminous infrared galaxy (ULIRG) by Hwang et al. (2007), thus supporting the presence of intense star-forming activity.

A search for the shortest timescale for a factor-of-two change in flux demonstrates intraday variability for 7/42 sources (see Table 9). There is some bias in this case, because not all *Swift*/UVOT observations were performed using all six filters. Two sources were extensively observed with almost all the six filters (J0324+3410 and J0948+0022) and displayed intraday variability at all wavelengths. This is in agreement of previous findings of extreme intraday optical variability reported by Liu et al. (2010), Itoh et al. (2013), Maune et al. (2013), and Paliya et al. (2013a). In particular, Itoh et al. (2013) report changes on timescale of minutes in the optical polarised flux of J0948+0022 on 2012 December 20, with a peak degree of polarisation of 36%.

Jiang et al. (2012) examined WISE data in a search for infrared variability among the 23 RLNLS1 of the Yuan et al. (2008) sample. They found three cases, also in the present sample: J0849+5108, J0948+0022, and J1505+0326. The first

two sources displayed intraday variability, while the latter showed significant flux changes over ~ 6 months. The remaining 20 RLNLS1s of the Yuan et al. (2008) sample did not show variability, most likely because of the weakness of the sources.

A more detailed analysis of the characteristics of the optical spectra (line, bumps, blue/red wings) will be presented elsewhere (Berton et al., in prep.).

4.4. Radio

About half of the sources (21/42) in the present sample were only detected at 1.4 GHz so it is not possible to determine a radio spectral index. In the remaining half of the cases, it was possible to estimate a spectral index between two frequencies below 8.4 GHz (between 1.4 and 5 GHz in 13/21 cases). About 28% of these sources (12/42) were detected at frequencies in the MHz range (74–843 MHz, see Table 9). In 4/42 cases (J0324+3410, J0849+5108, J0948+0022, and J1505+0326), spectral indices between 5–15 and 15–37 GHz are measured, because of the MW campaign of Effelsberg, Metsähovi, and RATAN-600 (Abdo et al. 2009a,b; Foschini et al. 2011a, 2012; Fuhrmann et al. 2011; Angelakis et al. 2012a,b). In only two cases (J0324+3140 and J0948+0022) are there detections up to 142 GHz at Pico Veleta. The results are summarised in Table 8. In 7/13 cases, the spectral indices α_r were inverted. Three of these cases were of γ -ray detected RLNLS1s. Two of these cases, J0324+3410 and J0849+5108 (two γ -ray sources), the average spectral index was inverted at higher frequencies. The weighted mean α_r was equal to 0.1 ± 0.3 (median 0.3).

Comparison with blazar samples reveals similar spectral indices. Abdo et al. (2010b) performed a linear regression of all the available data in the 1–100 GHz frequency range of 48 blazars of the *Fermi* LAT Bright AGN Sample (LBAS) and find an average value of $\alpha_{1-100\text{GHz}} = 0.03 \pm 0.23$. They found no differences between FSRQs and BL Lac objects. It is worth noting that the RLNLS1 J0948+0022 of the present sample is also in the LBAS list, but it is classified there as a low-synchrotron peak FSRQ. Abdo et al. (2010) find a radio spectral index of -0.645 .

Another useful comparison is with the jetted AGNs of the MOJAVE sample: Hovatta et al. (2014) analysed the radio data of 191 AGNs (133 FSRQs, 33 BL Lac objects, 21 radio galaxies, and 4 unknown-type AGNs) and calculated the spectral index between 8.1 and 15.4 GHz. Also in this case, there is virtually no difference between FSRQs and BL Lac objects, -0.22 and -0.19 , respectively.

Tornikoski et al. (2000) studied 47 Southern hemisphere sources, mostly FSRQs (38), plus 6 BL Lac objects and 3 radio galaxies (see also Ghirlanda et al. 2010 for a sample of blazars in the southern hemisphere) over a frequency range spanning 2.3 to 230 GHz. The spectral indices are almost flat below 8.4 GHz, with some differences between BL Lacs and high-polarisation quasars on one hand, and low-polarisation quasars on the other: while the latter have a somewhat steeper spectral index ($\alpha_{2.3-8.4\text{GHz}} \approx 0.05$), the former have an inverted index ($\alpha_{2.3-8.4\text{GHz}} \approx -0.13$). For all the sources, the spectral index becomes steeper at higher frequencies ($\alpha_{90-230\text{GHz}} \approx 0.79$). Nieppola et al. (2007) studied a large sample (398) of only BL Lac objects in the northern hemisphere and found average values of $\alpha_{5-37\text{GHz}} \approx -0.25$ and $\alpha_{37-90\text{GHz}} \approx 0.0$.

Our values are in agreement with these results: we find a rather flat or inverted spectrum extending from all the available frequencies (Table 8), as already found by Fuhrmann et al. (2011) and Angelakis et al. (2012a,b, 2015). However, we note that most of the radio observations refer to the first γ -ray

RLNLS1s detected, which were immediately monitored with MW campaigns, i.e. J0324+3410, J0849+5108, J0948+0022, and J1505+0326. All the other sources in the present sample have been poorly observed in the radio. There are, however, programs to increase the database on these sources at radio frequencies. The Metsähovi group is performing a multi-epoch variability program at 22 and 37 GHz on more than 150 radio-loud AGNs, including 38 RLNLS1s of the present sample (Lähtenmäki et al., in prep.). Richards et al. (2014) is performing a high-spatial resolution study on 15 RLNLS1s of the present sample with the VLBA at 5, 8, 15, and 24 GHz, including polarisation.

Analysis of archival VLBI radio maps provides some further information (see Sect. 3.2 and Table 7). Again, the earliest γ -ray detected RLNLS1s were the most heavily observed: J0324+3410, J0849+5108, and J0948+0022. The source J0324+3410 (7 observations at 15 GHz) displayed a small flare in 2011, with compactness – defined as the core to total flux density ratio – of order 0.6, which increased during the flare to >0.75 . Polarisation is almost stable during these observations, but during the flare there is a marginal change in the position angle of the electric vector (EVPA) with respect to the jet direction (from 87° to 77°). Also J0849+5108 showed an increase in the compactness (from 0.75 to 0.9) with flux density. Between 2013 January and July (MJD 56313–56481), there was a swing in the EVPA (from 168° to 24° , with an almost stable jet direction) that preceded a γ -ray outburst (Eggen et al. 2013), which happened also in J0948+0022 (Foschini et al. 2011a). The latter has been observed 17 times at 15 GHz, and many of these epochs were already studied by Foschini et al. (2011a). The present reanalysis basically confirms and extends the previous studies. It is worth noting that this source is very compact (0.975) as also indicated by previous comparison with single-dish observations (Abdo et al. 2009b; Foschini et al. 2011a, 2012). An exceptional radio core outburst on 2013 May 5 (MJD 56417), when J0948+0022 reached a core flux density of 0.862 Jy. This followed a strong outburst at γ -rays that occurred on 2013 January 1 (MJD 56293), as reported by D’Ammando et al. (2013c). During the same period there was also a swing of the EVPA, changing from $\sim 82^\circ$ on 2012 November 11 to $\sim 125^\circ$ on 2013 May 5. Moreover, Itoh et al. (2013) reported strong variability in optical polarisation in the same epoch (see Sect. 4.3).

The only source for which there are multifrequency observations is J1505+0326, from 2 to 43 GHz. A detailed analysis is reported by D’Ammando et al. (2013a). It is very compact (0.95–0.97 at 15 GHz during flares), at level comparable to J0948+0022. A flux density excess (~ 0.1 Jy) outside the core has been measured at 22 GHz in 2002. The EVPA–jet direction angle is quite unstable, changing from $\sim 61^\circ$ to -100° . We observed significant flux density increases only in the VLBI cores of all the observed sources (Table 8), suggesting that the location of the γ -ray emission should be very close to the central black hole.

Morphological studies of RLNLS1s have been published by Doi et al. (2006, 2007, 2011, 2012), Gu & Chen (2010), Giroletti et al. (2011), and Orienti et al. (2012). The emerging characteristics are (a) compact radio morphology, although there are kiloparsec scale structures in some cases (Doi et al. 2012); (b) high-brightness temperature of the core feature, indicating non-thermal processes; (c) flat or inverted spectra (although the samples included also steep spectrum radio sources, which are excluded from the present work); and (d) possible links with compact steep spectrum (CSS) and gigahertz peaked spectrum (GPS) radio sources, as also suggested by Oshlack et al. (2001)

and Gallo et al. (2006) for J2007–4434, Komossa et al. (2006a), Yuan et al. (2008), and more recently by Caccianiga et al. (2014). Doi et al. (2012) found that the detection of extended emission is lower than expected from broad-line Seyferts and they suggest it could be due to the lower kinetic power of jets in low-mass AGNs, rather than the young age of the source. Interestingly, also the radio core of RQ-NLS1s and Seyferts display non-thermal characteristics, suggesting some link with jets (Giroletti & Panessa 2009; Doi et al. 2013).

Angelakis et al. (2015) have studied the variability of four RL-NLS1s detected at γ rays (J0324+3410, J0849+5108, J0948+0022, J1505+0326) at different radio frequencies. Brightness temperature measurements indicate minimum Doppler factors from 1.3 for J0324+3410 to 4.2 in the case of J1505+0326.

The search for short variability resulted in only one case of variability on timescales of days: we measure an upper limit of 2.6 days for J0948+0022 at 37 GHz. In the other cases, we find variations on timescales of about one month, but this is likely due to the one-month sampling rate of Effelsberg observations. Metsähovi (37 GHz) observed at a more intense sampling rate during some MW campaigns (e.g. Abdo et al. 2009b; Foschini et al. 2011a, 2012). Nieppola et al. (2007) reported variability on timescales of hours for some BL Lac objects, for example, ~ 8 h for S5 0716 + 71, ~ 1 hour for AO 0235+164, and ~ 6 h for OJ 287. In the case of RL-NLS1s, clearly higher sampling rate MW campaigns are required.

5. Estimates of masses and accretion luminosities

The masses of the central black holes are given by

$$M = f \left(\frac{R_{\text{BLR}} \sigma_{\text{line}}^2}{G} \right), \quad (2)$$

where R_{BLR} is the size of the broad-line region (BLR) measured by reverberation or estimated from scaling relations, σ_{line} is the line dispersion (or second moment of the line profile), G is the gravitational constant, and f is a dimensionless scale factor of order unity (Peterson et al. 2004). We used the line dispersion, because it is less affected by inclination, Eddington ratio, and line profile (Peterson et al. 2004; Collin et al. 2006). We estimate the BLR radius by using the relationship between the luminosity of the $H\beta$ line and the radius of the BLR (R_{BLR}) from the relationship of Greene et al. (2010),

$$\log \left[\frac{R_{\text{BLR}}}{10 \text{ light days}} \right] = 0.85 + 0.53 \log \left[\frac{L(H\beta)}{10^{43} \text{ erg s}^{-1}} \right]. \quad (3)$$

Following Collin et al. (2006), we adopt $f = 3.85$.

The size of the BLR gives also the luminosity of the accretion disk ($R_{\text{BLR}} \propto L_{\text{disk}}^{1/2}$, e.g. Ghisellini & Tavecchio 2009; see also Bentz et al. 2013), which in turn has been normalised to the Eddington value

$$L_{\text{Edd}} = 1.3 \times 10^{38} \left(\frac{M}{M_{\odot}} \right) \left[\text{erg s}^{-1} \right]. \quad (4)$$

By using $L(H\beta)$ – instead of the continuum at 5100 Å (or another wavelength), which is more conventional and generally more accurate – to estimate the size of the BLR and the accretion disk luminosity, we avoid the problem of contamination of the flux by either the jet or the host galaxy.

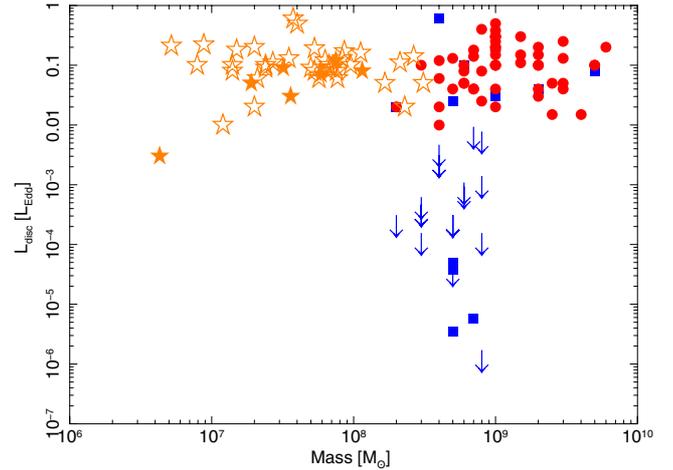


Fig. 4. Accretion disk luminosity [Eddington units] vs. mass of the central black hole [M_{\odot}]. The orange stars are the RL-NLS1s of the present sample (see Table 2) and filled orange stars indicate those detected at γ rays; the red circles are the FSRQs, and the blue squares are the BL Lac objects (blue arrows indicates upper limits in the accretion luminosity) from Ghisellini et al. (2010). We noted some BL Lacs with strong accretion disk, in the region occupied by FSRQs: these are the so-called intruders (Ghisellini et al. 2011; Giommi et al. 2012).

The results are displayed in Table 2. In three cases (3/42), no optical spectra were found. Therefore, we estimated the line dispersion from the value of FWHM found in literature by using the ratio $FWHM/\sigma_{\text{line}} = 1.07$, which is the average over the known values of the present sample (39/42). This value is consistent with what expected from NLS1s (cf. Peterson 2011). From the available optical magnitudes near 5100 Å we estimated the disk luminosity and the size of the BLR and then used Eq. (2) to estimate the mass. We note that these sources have Eddington ratios slightly greater than the others of the sample: this can be understood because with the photometry it is not possible to disentangle the contribution of the disk from that of the jet. We note that our values are in agreement with the results available in the vast majority of literature on RL-NLS1s (e.g. Komossa et al. 2006; Whalen et al. 2006; Yuan et al. 2006) and on NLS1s in general (e.g. Peterson 2011).

A comparison of these data with the corresponding data for the blazar sample is shown in Fig. 4. It is evident that RL-NLS1s occupy a unique parameter space among AGNs with relativistic jets that corresponds to lower masses and high Eddington rates. It is worth noting one outlier, J2007–4434, has a low Eddington rate ($0.003 L_{\text{Edd}}$): this was also one of the RL-NLS1s whose nature is questionable on account of its weak Fe II emission (Gallo et al. 2006; Komossa et al. 2006). There is apparently an unoccupied area of parameter space corresponding to low black hole masses and low Eddington ratios. It is not possible to say whether this is real or a selection effect. Possible candidates to occupy this region are low-luminosity AGNs (e.g. M 81, Alberdi et al. 2013), although there is debate about the nature of their radio emission (see Paragi et al. 2013). Moreover, no low-luminosity AGN has yet been detected in high-energy γ rays.

6. Monochromatic luminosities

Another comparison with blazars can be done via νL_{ν} - νL_{ν} plots in Fig. 5. Starting from the available data, we normalised the fluxes to four reference frequencies or wavelengths or energies: 15 GHz for radio observations, 203 nm for ultraviolet

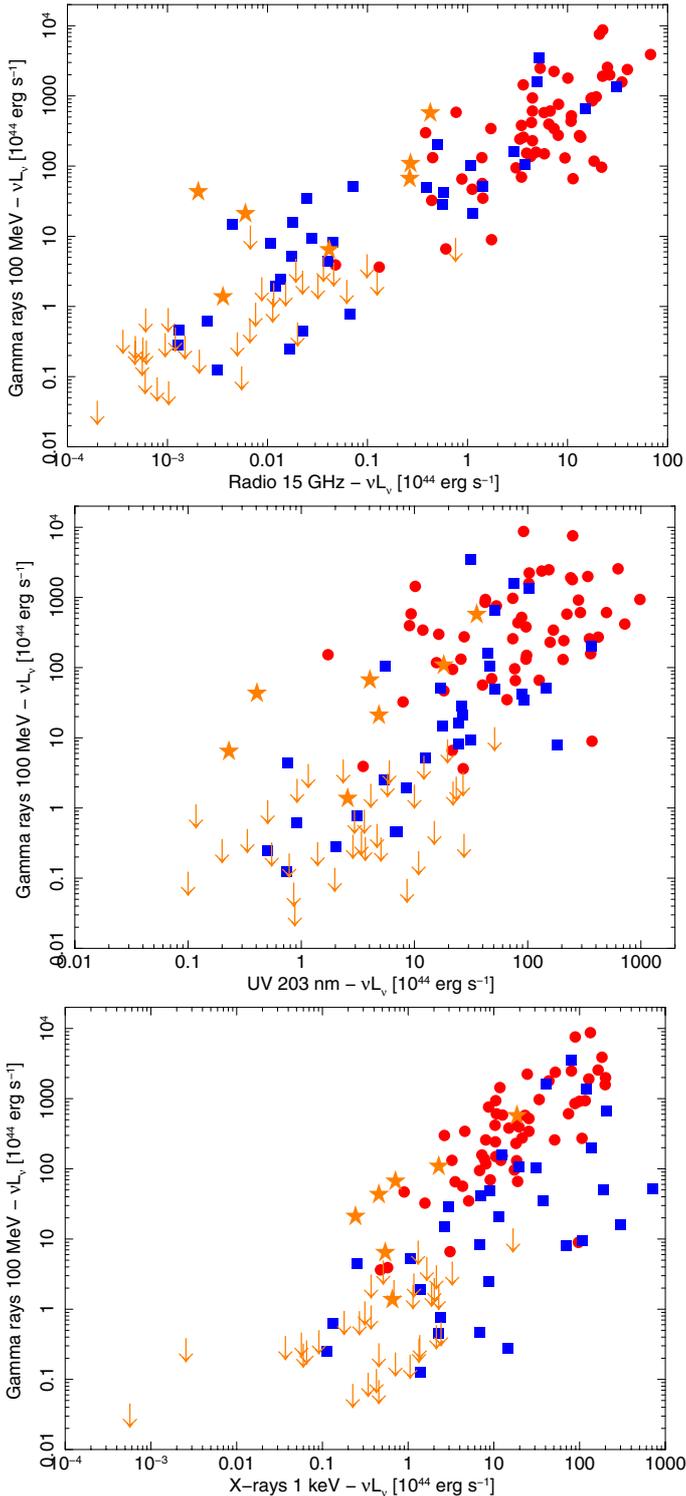


Fig. 5. Gamma-ray luminosity at 100 MeV compared with radio luminosity at 15 GHz (*top panel*), ultraviolet luminosity at 203 nm (*middle panel*), X-ray luminosity at 1 keV (*bottom panel*). The orange stars are the RLNLS1s of the present sample detected in γ rays, while upper limits are reported for the others (grey arrows); the red circles are the FSRQs and the blue squares are the BL Lac objects.

wavelengths, 1 keV for X-rays, and 100 MeV for γ -rays. While for most blazars, radio observations at 15 GHz were available from the MOJAVE project, the same was not true for RLNLS1s. For half of the RLNLS1s (21/42), there were radio data at 1.4 GHz either from NVSS or FIRST surveys. In some cases,

there were also data at 5, 8.4, 17, or 20 GHz (the two latter frequencies are used in the Southern hemisphere). We extrapolated the 15 GHz flux by using the average radio spectral indices in Table 8.

The situation is slightly better at ultraviolet wavelengths, because of the availability of *Swift*/UVOT observations, many of them specifically requested for this survey. For those sources with incomplete data, we used the bluest photometric data available and corrected by using the average UV spectral index in Table 8. We adopted the average spectral indices also to normalise the integrated fluxes or upper limits in the 0.3–10 keV and 0.1–100 GeV bands.

The monochromatic fluxes were then K -corrected by

$$S_{\nu, \text{rest}} = S_{\nu} (1+z)^{\alpha_{\nu}-1}, \quad (5)$$

where $S_{\nu, \text{rest}}$ is the rest-frame monochromatic flux at the frequency ν , S_{ν} is the observed monochromatic flux at frequency ν , z is the redshift, and α_{ν} is the spectral index at the frequency ν . The corrected monochromatic fluxes were then converted into luminosities. The results are displayed in Fig. 5.

Figure 5 shows that RLNLS1s are the low-luminosity tail of FSRQs, as already noted by Foschini et al. (2013). While at radio and ultraviolet frequencies RLNLS1s share the same region of BL Lac objects, the two populations diverge from each other at 1 keV, where BL Lac objects move to greater X-ray luminosities, indicating a different origin of the emission (synchrotron for BL Lacs, disk corona or inverse-Compton for RLNLS1s).

We noted one possible outlier in the radio- γ panel: J1102+2239 was detected at γ -ray flux above what is expected from the trend of the other sources. There could be several explanations: the γ -ray activity could be limited to a small time interval, the radio measurements, extrapolated from 1.4 GHz measurements from FIRST and NVSS, were likely done during periods of low activity of the sources, or it could even be an indication of some artefact in the γ -ray detection. Further studies could solve the conundrum.

There are also two sources with very low X-ray fluxes in the X-ray/ γ -ray panel, J0100–0200 and J0706+3901. In both cases, the X-ray flux was measured by *Chandra* in 2003 and there were no simultaneous data at other wavelengths.

We stress the difference between RLNLS1s and BL Lac objects. Figure 5 shows the observed luminosities at different frequencies: RLNLS1s and BL Lacs occupy similar regions and generally overlap at radio and UV frequencies. However, while BL Lac objects have low power and masses comparable to those of FSRQs, RLNLS1s have low power and lower masses (see Fig. 4). Indeed, when normalised for the mass of the central black hole, the jet power of RLNLS1s and FSRQs are of the same order of magnitude, as shown by Foschini (2014) and references therein. It is worth stressing that the normalisation is not linear, but it is necessary to divide the jet power by $M^{1.4}$, according to the theory developed by Heinz & Sunyaev (2003) and confirmed by Foschini (2011b, 2012b,c, 2014).

7. Spectral energy distribution

Figures 8–13 display the observed SEDs of all RLNLS1 in the present sample, assembled from data extracted from observations at different epochs and archives, as discussed in Sect. 3. The most complete SEDs are mostly those of the RLNLS1s significantly detected at gamma rays, i.e. J0324+3410, J0849+5108, J0948+0022, J1505+0326, J2007–4434. The

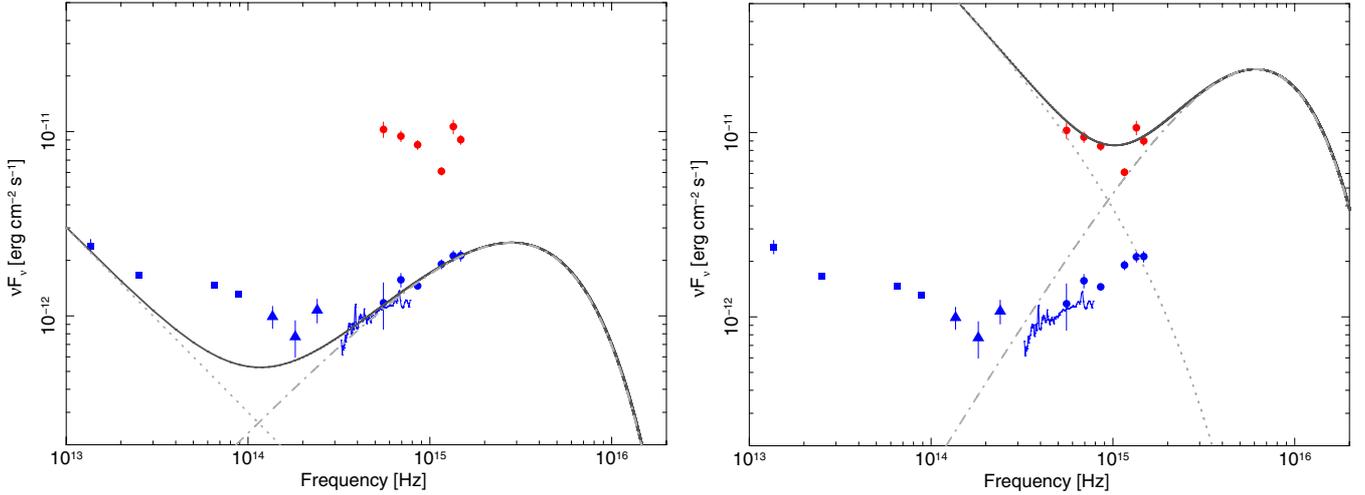


Fig. 6. Zoom of the SED of J0948+0022 in the infrared-to-ultraviolet range. Data are from: WISE (filled squares), 2MASS (filled triangles), SDSS (continuous line), *Swift*/UVOT (filled circles). *Left panel*: blue refer to lowest observed activity state (LS, 2009 May 15); *right panel*: red to highest activity state (HS, 2012 December 30). The grey dot-dashed line represents a model of standard accretion disk as expected in the case of J0948+0022 ($M = 7.5 \times 10^7 M_{\odot}$); the grey dotted line represents the synchrotron emission; the continuous grey line is the sum of the two models.

modelling of these SEDs already has been presented and discussed in other papers (e.g. Abdo et al. 2009b,c; Foschini et al. 2011a, 2012; D’Ammando et al. 2012, 2013a,b; Paliya et al. 2014).

There are also a few more cases (e.g., J0814+5609, J1047+4725, J1548+3511, J1629+4007) with fairly good sampling because of previous specific interest. For example, J1629+4007 was long observed because it was thought to be an example of a high-frequency peaked FSRQ (Padovani et al. 2002; Falcone et al. 2004). It is evident from the SED (Fig. 15) that the strong X-ray emission is not due to synchrotron radiation, but rather to the disk corona (see also Maraschi et al. 2008). In other cases, it seems simply that the source fell into the field of view of other targets. We noted a strong change in radio flux density at 5 GHz in J1047+4725: early observations performed in 1987 with the Green Bank 91 m telescope (~ 3.5 angular resolution) found a flux density of ~ 0.4 Jy (Becker et al. 1991; Gregory & Condon 1991), while an observation with the VLA at 8.4 GHz in 1990 (0.2 angular resolution) measured a flux density of ~ 0.3 Jy (Patnaik et al. 1992). A VLBA observation at 5 GHz with milliarcsecond resolution in 2006 resulted in a flux density of ~ 33 mJy (Helmboldt et al. 2007). One explanation for this could be the presence of significant extended emission, which is integrated in the low angular resolution of Green Bank 91 m and VLA telescopes, while is resolved in the milliarcsecond images of VLBA.

Other sources displayed extreme variability, specifically at optical wavelengths: for example, the SDSS optical spectrum of J0849+5108 (observed in 2000) is about one order of magnitude lower than the optical observations made with *Swift* after the detection at γ -rays by *Fermi* (2011–2013). Similar cases are J1159+2838 and J2007–4434, while the optical spectrum of J0953+2836 is about two orders of magnitude brighter than in the *Swift* observations. Spectral changes at optical frequencies, due to the jet activity, are also observed. Just as an example, we focus on the infrared-to-ultraviolet band of J0948+0022, which is the best sampled source, being the first to be detected at γ -rays. Figure 6 displays the two extreme states from the available data: the lowest activity state (LS, 2009 May 15, see also Abdo et al. 2009b; Foschini et al. 2012) and the highest state

(HS, 2012 December 30). In both cases, we model the synchrotron emission (dotted grey line) with a power-law model with an exponential cutoff. The disk emission (dashed grey line) is the standard Shakura-Sunyaev model as expected from a black hole of $M = 7.5 \times 10^7 M_{\odot}$ (see Table 2). Previous modelling (Abdo et al. 2009b; Foschini et al. 2012) supposed constant disk luminosity equal to $L_{\text{disk}} \sim 0.4 L_{\text{Edd}}$, as measured by fitting the optical/UV emission with a standard Shakura-Sunyaev disk. In the present work, we obtained from the emission lines a value of $L_{\text{disk}} \sim 0.12 L_{\text{Edd}}$. The difference could be due to a contamination of the jet emission in the optical/UV photometry fit, which is removed by using the emission lines. However, since the optical spectrum of SDSS was observed in 2000 and the MW campaign used for the SED modelling were obtained in 2008–2011, it is also possible that the Eddington ratio really changed.

Although the lowest flux points were not simultaneous, since many of them are smoothly connected, it is reasonable to assume that they refer to a common state of low jet activity. Therefore, we modelled them together as low activity state (LS): in this case, the expected peak of the emission from the accretion disk at 12% of the Eddington luminosity is at $\sim 3.0 \times 10^{15}$ Hz. For the high state (HS), we have just one *Swift*/UVOT observation and we adopted a standard disk at 40% of the Eddington value, peaking at $\sim 4.1 \times 10^{15}$ Hz. The cutoff of the synchrotron emission was set to 8×10^{14} Hz in the LS and increased to 1.5×10^{15} Hz in the HS. We underscore that this is just another option, in addition – and not in contrast – to the previous model. As the expected peak of the disk emission in the far UV, outside the range of our observations, we cannot clearly distinguish between different possibilities. We also note that at infrared wavelengths, there is an excess that is likely attributable to the dusty torus and/or the host galaxy, as suggested also by the WISE colours (see Sect. 4.3).

Changes in the activity of a relativistic jet, as for J0948+0022, could explain the differences in the spectral slopes of J0849+5108, J1159+2838, J2007–4434. All these sources have an optical spectrum with a slope different from that derived from the optical/UV photometry. Moreover, some sources show optical/UV slopes decreasing with increasing frequencies (e.g. J0804+3853, J0937+3615, J1031+4234, J1038+4227,

Table 3. Estimated jet power from radio core measurements at 15 GHz according to the relationships in Foschini (2014).

Source	$\log P_{\text{rad}}$	$\log P_{\text{kin}}$	$\log P_{\text{tot}}$
J0100 – 0200	42.88	43.06	43.28
J0134 – 4258	43.28	43.53	43.72
J0324 + 3410	43.17	43.40	43.60
J0706 + 3901	42.22	42.27	42.55
J0713 + 3820	42.67	42.81	43.05
J0744 + 5149	43.55	43.86	44.03
J0804 + 3853	42.59	42.71	42.96
J0814 + 5609	43.88	44.26	44.41
J0849 + 5108	44.58	45.09	45.21
J0902 + 0443	44.33	44.79	44.92
J0937 + 3615	42.56	42.67	42.92
J0945 + 1915	43.37	43.64	43.83
J0948 + 0022	44.72	45.27	45.38
J0953 + 2836	44.25	44.70	44.83
J1031 + 4234	43.54	43.85	44.02
J1037 + 0036	44.00	44.40	44.55
J1038 + 4227	42.57	42.68	42.93
J1047 + 4725	44.92	45.50	45.60
J1048 + 2222	42.59	42.71	42.95
J1102 + 2239	42.98	43.18	43.40
J1110 + 3653	43.92	44.31	44.46
J1138 + 3653	43.41	43.70	43.88
J1146 + 3236	43.64	43.97	44.14
J1159 + 2838	42.51	42.61	42.86
J1227 + 3214	42.58	42.70	42.95
J1238 + 3942	43.72	44.06	44.22
J1246 + 0238	43.34	43.61	43.79
J1333 + 4141	42.51	42.61	42.87
J1346 + 3121	42.42	42.50	42.76
J1348 + 2622	43.38	43.65	43.84
J1358 + 2658	42.76	42.91	43.14
J1421 + 2824	43.77	44.12	44.28
J1505 + 0326	44.57	45.08	45.20
J1548 + 3511	44.10	44.52	44.66
J1612 + 4219	42.74	42.88	43.12
J1629 + 4007	43.73	44.07	44.24
J1633 + 4718	42.76	42.91	43.14
J1634 + 4809	43.46	43.75	43.93
J1644 + 2619	43.31	43.57	43.76
J1709 + 2348	42.42	42.58	42.81
J2007 – 4434	43.96	44.36	44.50
J2021 – 2235	42.99	43.19	43.40

Notes. Columns display the logarithm of the radiative, kinetic (protons, electrons, magnetic field), and total jet powers [erg s^{-1}].

J1102+2239, J1138+365, J1227+3214), while there were other cases with the opposite trend (e.g. J0134–4258, J0324+3410, J0814+5609, J1348+2622, J1548+3511, J1629+400). These also have a flat X-ray slope, with some evidence of a soft-excess. An inspection of their corresponding central black hole masses and Eddington ratios did not reveal any trend. On the basis of the J0948+0022 behaviour, we favour the interpretation of the same central engine observed in a different combination of jet–disk states.

There are some cases with only a few data points, so that it is not possible to draw useful inferences (e.g. J0100–0200, J0706+3901, J1333+4141, J1346+3121, J1358+2658, J1612+4219, J1709+2348). While we have added significantly to the MW database for many of these objects, the radio observations remained limited to only 1.4 GHz. It is therefore desirable that future observations focus on radio frequencies (e.g. Richards et al. 2014; Lähteenmäki et al., in prep.).

8. Jet power

To estimate the jet power, we adopted the relationships based on the radio core measurements at 15 GHz by Foschini (2014),

$$\log P_{\text{jet,radiative}} = (12 \pm 2) + (0.75 \pm 0.04) \log L_{\text{radio,core}} \quad (6)$$

and

$$\log P_{\text{jet,kinetic}} = (6 \pm 2) + (0.90 \pm 0.04) \log L_{\text{radio,core}}. \quad (7)$$

From the values calculated in the Sect. 6, we derived the radiative, kinetic (protons, electrons, magnetic field), and the total jet power for each source. The results are given in Table 3.

In some cases, it is possible to test the present results with calculations performed by modelling the SED, with the caveat that we are comparing different epochs of strongly variable sources. For example, J0948+0022 – the first RLNLS1 to be detected at gamma rays – had a radiative jet power of $\log P_{\text{radiative}} = 45.5$, while the kinetic part was estimated as $\log P_{\text{kinetic}} = 46.9$ (Abdo et al. 2009a). During the 2009 MW campaign, these values ranged from 44.9 to 45.54 for the radiative power, and from 45.67 to 46.2 for the kinetic power (Abdo et al. 2009b). During more than three years of monitoring, $\log P_{\text{rad}}$ spanned the interval 44.55–45.97, while $\log P_{\text{kinetic}}$ was in the range 46.19–47.61 (Foschini et al. 2012). The present estimate (Table 3) is an average of several measurements done directly at 15 GHz (mostly by the MOJAVE project and Effelsberg), and is reasonably consistent with the previously published values (see also Angelakis et al. 2015). The greater values were recorded during the exceptional 2010 outburst, when J0948+0022 reached an observed luminosity of about $10^{48} \text{ erg s}^{-1}$ (Foschini et al. 2011a).

In other cases:

- J0324+3410: $\log P_{\text{rad}} = 42.8$, $\log P_{\text{kin}} = 44.3$ (averaged over one year, Abdo et al. 2009c), and $\log P_{\text{rad}} = 41.29$ –41.74, $\log P_{\text{kin}} = 44.06$ –45.14 (different states over five years monitoring, Paliya et al. 2014).
- J0849+5108: $\log P_{\text{rad}} = 45.6$ (peak during an outburst), $\log P_{\text{kin}} = 45.3$ ¹⁰ (D’Ammando et al. 2012).
- J1505+0326: $\log P_{\text{rad}} = 44.0$, $\log P_{\text{kin}} = 46.2$ (averaged over one year, Abdo et al. 2009c).
- J2007–4434: $\log P_{\text{rad}} = 42.9$, $\log P_{\text{kin}} = 44.1$ (averaged over one year, Abdo et al. 2009c).

A comparison with the jet power of FSRQs and BL Lac objects (Fig. 7) shows that RLNLS1s have values comparable to BL Lac objects but lower than FSRQs. The mean values are $\log P_{\text{rad}} = 43.35$ and $\log P_{\text{kin}} = 43.62$ for RLNLS1s, $\log P_{\text{rad}} = 45.49$ and $\log P_{\text{kin}} = 46.78$ for FSRQs, and $\log P_{\text{rad}} = 44.14$ and $\log P_{\text{kin}} = 45.01$ in the case of BL Lac objects. Taking into account a mean value for the masses of the central black holes of the three populations ($M_{\text{RLNLS1}} = 6.8 \times 10^7 M_{\odot}$, $M_{\text{FSRQs}} = 1.5 \times 10^9 M_{\odot}$, and $M_{\text{BLLacs}} = 7.2 \times 10^8 M_{\odot}$) and renormalizing by $M^{1.4}$, we obtained $\log P_{\text{rad}} = 32.38$ and $\log P_{\text{kin}} = 32.65$ for RLNLS1s, $\log P_{\text{rad}} = 32.64$ and $\log P_{\text{kin}} = 33.93$ for FSRQs, and $\log P_{\text{rad}} = 31.74$ and $\log P_{\text{kin}} = 32.61$ in the case of BL Lac objects. Thus, the normalised jet power is almost the same for all the three types of AGNs, as expected (see Sect. 6), and it is also consistent with the jets from Galactic binaries (Foschini 2014).

¹⁰ Electrons and magnetic field only.

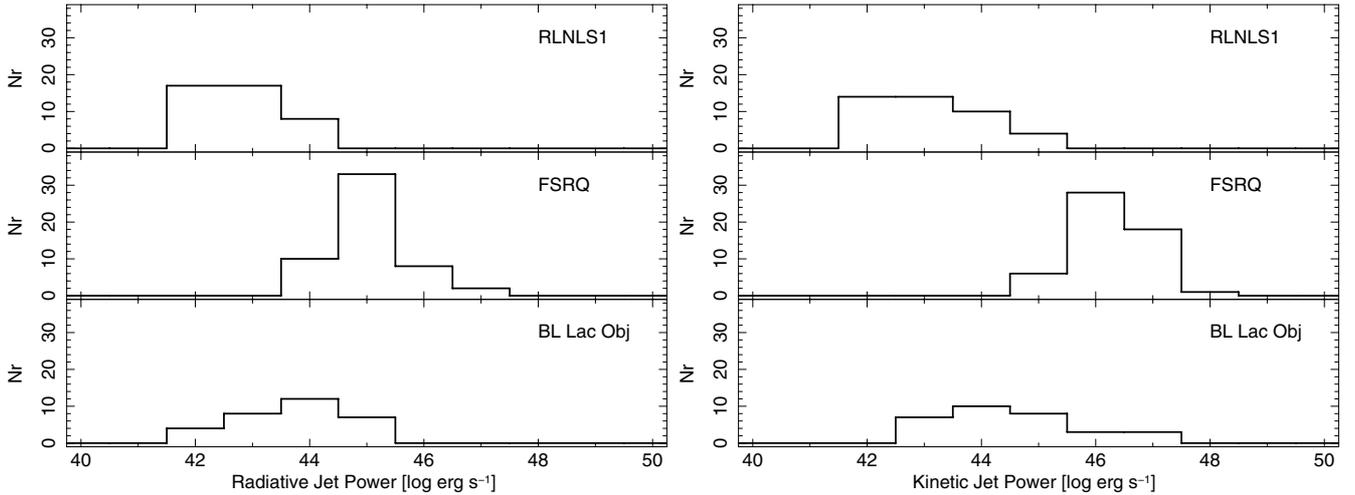


Fig. 7. Jet power distribution: (*left panel*) radiative, (*right panel*) kinetic. RLNLS1s data are from the present work (Table 3), while values for FSRQs and BL Lac objects are from Ghisellini et al. (2010).

9. Discussion

Since the discovery of NLS1s, there has been debate as to whether they are an intrinsically separate AGN class, or simply the low-mass tail of the distribution of Seyferts (Osterbrock & Pogge 1985). Many authors favoured the latter hypothesis (e.g. Grupe 2000; Mathur 2000; Botte et al. 2004). The same question has been proposed in the case of RLNLS1s (Yuan et al. 2008). The first studies following the detections at γ rays suggested a simple mass difference (Abdo et al. 2009a,c; Foschini 2011a, 2012a; Foschini et al. 2011a, 2013). The unification of relativistic jets provided further support for this point of this view (Foschini 2011b, 2012b,c, 2014). On the basis of what we have found in this survey, with more sources and data, we can confirm that, although RLNLS1s show some peculiar observational differences with respect to the other radio-loud AGNs (the optical spectrum and the possible starburst activity), the physical characteristics inferred from the data (mass of the central black hole, Eddington ratio, spectrum, jet power) favour the hypothesis that RLNLS1s are the low-mass tail of AGNs with jets. This is one more point favouring the Livio (1997) conjecture, according to which the jet engine is the same, but the observational features are different, depending on a number of variables, such as the mass of the central accreting body, the accretion flow, and the local environment.

In the case of RLNLS1s, the relatively lower mass of the central black hole implies variability on very short timescales, much smaller than expected only from Doppler boosting, which is exactly what is seen when the observational coverage allows it. It is known that the power spectral densities of AGNs show a break timescale, t_b , separating long-term timescales from the shorter ones (McHardy et al. 2006). There are some relationships linking t_b with the mass of the central black hole, the bolometric luminosity, or the FWHM of the H β (McHardy et al. 2006; González-Martín & Vaughan 2012). By taking as representative values the averages of the selected quantities, it is possible to estimate t_b , which is expected to be around minutes to hours for RLNLS1s, and hours to a few days for blazars. Indeed, hour timescales at high energies are exceptional events for blazars (e.g. Foschini et al. 2011b,c), but are quite common for RLNLS1s as there are sufficient statistics to allow a meaningful detection (Table 9). As stated in Sect. 4.2, the claim of minute

timescale X-ray variability requires further detailed study, but it is worth noting the 2–3 min timescale variability in the optical polarisation reported by Itoh et al. (2013).

We have observed not only flux variability, but also spectral changes, suggesting the interplay of jet and disk components (see the case of J0948+0022 in Sect. 7). At a first look, the SEDs suggest two different classes of RLNLS1s, depending on the slope of the optical/UV spectra. However, the spectral variability of some sources (e.g. J0849+5108, J0945+1915, J0948+0022, J1159+2838, J2007–4434) simply indicates that we are observing different states of activity of the same central engine. Indeed, the two classes do not show any difference in the mass, disk, and jet parameters.

The lower mass of the central black holes in RLNLS1s has an important implication: the observed jet luminosity is lower than that of quasars, but comparable to that of BL Lac objects. Therefore, one could wonder why the RLNLS1s are more difficult to discover than BL Lacs? The latter are generally more luminous at X-rays than RLNLS1s because the synchrotron radiation peaks in the UV/X-rays (see Fig. 5, bottom panel), and indeed, BL Lac objects are more easily found in X-ray surveys (Padovani & Giommi 1995). At γ rays, *Fermi*/LAT discovered many BL Lac objects because the instrumental characteristics of LAT favour hard sources at low fluxes: this made it easier to detect BL Lacs ($\alpha_\gamma < 1$), but not RLNLS1s ($\alpha_\gamma > 1$) (see Sect. 4.1). At radio frequencies, both RLNLS1s and BL Lac objects are weak (see Fig. 5, top panel). However, Giroletti et al. (2012) noted that BL Lacs have extended radio emission, which is almost missing in RLNLS1s (e.g., Doi et al. 2012). One possible explanation, advanced by Doi et al. (2012), is that in the case of RLNLS1s, the jet has low kinetic power because of the low mass and because it has to propagate in a gas-rich environment, while in BL Lacs the jet power is slightly greater and develops in a more rarified medium. Another possibility is to invoke the young age of RLNLS1s (Mathur 2000; Mathur et al. 2012) and, indeed, many authors made the hypothesis of a link with GPS/CSS sources, which in turn are believed to be very young radio galaxies (Oshlack et al. 2001; Komossa et al. 2006a; Gallo et al. 2006; Yuan et al. 2008; Caccianiga et al. 2014). Yet another option has been proposed by Gu & Chen (2010): the jet activity could be intermittent, as observed in other Seyferts (e.g. Brunthaler et al. 2005; Mundell et al. 2009). Therefore, as the

technological improvement of radio surveys allows better monitoring of these sources (e.g. Square Kilometer Array, SKA), the rate of detection should increase.

The intermittent jet should not be confused with the outburst/flare activity as observed in blazars. In the case of RLNLS1s, the periods of activity/inactivity might be separated by dramatic changes in flux. Indeed, in addition to the episodes of strong variability already described (see Sect. 7), we also note two sources where the X-ray flux was three-to-four orders of magnitude lower than the optical/UV emission (J0100–0200, J0706+3901). In other cases, although the SED displayed the double-humped shape typical of a domination of a relativistic jet, the lack of γ -ray detection (no new detection was reported to date) set very stringent constraints (e.g. J0814+5609, J1031+4234, J1421+2824, J1629+4007). This can be compared with the lowest-known state of the BL Lac object PKS 2155–304 ($z = 0.116$) where the changes in the X-ray flux were of about an order of magnitude and there was a shift of the synchrotron peak at lower frequencies (Foschini et al. 2008). This indicates a jet with a continuous background of emitted radiation, with superimposed outbursts and flares, as new blobs are ejected. The more dramatic changes of three-to-four orders of magnitudes observed in RLNLS1s suggests that the central engine changes its level of activity significantly: not only the jet, but also the corona seems to be strongly reduced.

Czerny et al. (2009), supported by Wu (2009), proposed a radiative instability in the accretion disk to explain the intermittent activity in young radio sources. RLNLS1s have all accretion luminosities sufficiently high to be in the radiation-pressure dominated regime (Moderski & Sikora 1996; Ghosh & Abramowicz 1997; see Foschini 2011b for the application to RLNLS1s), where Czerny’s theory applies. The timescale of the active phase in the case of low-mass AGNs, such as NLS1s, could be very small, of the order of tens-to-hundreds of years (Czerny et al. 2009). Therefore, the low kinetic power of the jet due to the low mass of the central black hole, the short periods of activity, and a frustrating nearby environment rich in interstellar gas and photons, are the sufficient ingredients to explain the lack of extended radio relics. As suggested by Doi et al. (2012), such structures might appear only in the sources with greatest black hole masses, which in turn might be in the final stages of their cosmological evolution before changing into broad-line AGNs.

Another possibility is the aborted jet model proposed by Ghisellini et al. (2004), which in turn could also explain the difference between radio-loud and radio-quiet AGNs. In this case, the jet has insufficient power to escape from the central black hole and falls back. The spectral characteristics in the X-ray band could be an index generally steeper than usual for Seyferts (that is $\alpha_X \sim 1$), significant equivalent-width fluorescent iron lines, and a steeper-when-brighter behaviour of the light curves. J0324+3410 might be a good candidate, also because it is the only one with a detected Fe $K\alpha$ line. However, the X-ray flux and spectral index values (Table 5) do not reveal any significant trend. We note that high-flux periods have both harder and steeper indices. We can speculate that a jet might sometimes be aborted (steeper when brighter) or launched (harder when brighter). The rather obvious question is then what determines one or the other?

10. Conclusions

We have presented a survey of 42 RLNLS1s observed from radio to γ rays, the largest MW sample to date of this type of source. In addition to previously published data, we present here new

analyses of data obtained with *Swift* and *XMM-Newton* specifically to address these sources. The main results of the analyses are:

- γ rays: 7/42 sources (17%) were detected at high-energy γ rays. The average spectral index is $\alpha_\gamma \sim 1.6$, consistent with that of FSRQs. Intraday variability has been reported in three sources.
- X-rays: we detected 38/42 sources (90%), with an average spectral index $\alpha_X \sim 1.0$ and median 0.8. We also detected variability on timescales of hours in 6 sources.
- Intraday variability was observed also at ultraviolet/optical wavelengths in those few sources which were targets of MW campaigns. Dramatic changes both in fluxes and spectra were also observed when comparing observations on timescales of years. Infrared colours indicate that RLNLS1s are basically on the line expected from synchrotron emission, but with a significant spread towards the starburst region.
- We observed in some sources changes of the EVPA corresponding to γ ray activity. We detected significant changes of radio flux density only in the VLBI-cores, suggesting that the emission of γ rays should occur close to the central black hole.
- Comparison of monochromatic luminosities at 15 GHz, 203 nm, 1 keV, and 100 MeV with a sample of blazars (FSRQs, and BL Lac objects) suggest that RLNLS1s are the low-power tail of the quasar distribution.
- Some SEDs confirm the dramatic variability already apparent from the single band analysis. We modelled one case (J0948+0022) to show how the observed spectral variability can be interpreted as the interplay of the jet and accretion disk emission.

The radio coverage are still deficient, but some programs are ongoing (Richards et al. 2014; Angelakis et al. 2015; Lähteenmäki et al., in prep.).

The main results calculated from the data are:

- The estimated masses of the central black holes ($10^{6-8} M_\odot$) and Eddington ratios ($0.01-0.49 L_{\text{Edd}}$) are in the range typical of NLS1s, with one outlier, J2007–4434, at $0.003 L_{\text{Edd}}$. The masses are lower than those of blazars ($10^{8-10} M_\odot$), indicating that we are studying a new different regime of the mass-accretion parameter space.
- The calculated jet powers ($10^{42.6-45.6} \text{ erg s}^{-1}$) are generally lower than those of FSRQs and partially overlapping, but still slightly lower than those of BL Lac objects. Once normalised by the mass of the central black holes, the jet powers of the three populations are consistent with each other, indicating the scalability of the jet.

The inferences that can be drawn from this study are that, despite the observational differences, the central engine of RLNLS1s is quite similar to that of blazars, as indicated by the scalability of the jet emission. The difficulties in finding this type of source might be due to the low observed power and an intermittent activity of the jet. Large monitoring programs with high-performance instruments (e.g. SKA) should allow us to greatly improve our understanding of these sources, which will lead to a better understanding of the more general issue of the physics of relativistic jets and how they are generated.

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Table 4. Gamma-ray spectral characteristics.

Name	Flux	Photon index	TS	Time period	Reference
J0324 + 3410	6.0 ± 0.7	2.87 ± 0.09	164	2008 Aug.–2011 Feb.	Foschini (2011a)
J0849 + 5108	0.51 ± 0.15	2.0 ± 0.1	52	2008 Aug.–2011 Feb.	Foschini (2011a)
	2.6 ± 0.2	2.18 ± 0.05	658	2008 Aug.–2012 Aug.	D’Ammando et al. (2013d)
J0948 + 0022	13.7 ± 0.7	2.85 ± 0.04	1081	2008 Aug.–2011 Feb.	Foschini (2011a)
	13.6 ± 0.3	2.67 ± 0.03	2015	2008 Aug.–2011 Dec.	Foschini et al. (2012)
J1102 + 2239	2.0 ± 0.6	3.1 ± 0.2	32	2008 Aug.–2011 Feb.	Foschini (2011a)
J1246 + 0238	1.7 ± 0.7	3.1 ± 0.3	15	2008 Aug.–2011 Feb.	Foschini (2011a)
J1505 + 0326	7.0 ± 0.6	2.71 ± 0.07	411	2008 Aug.–2011 Feb.	Foschini (2011a)
	5.1 ± 0.4	2.67 ± 0.06	419	2008 Aug.–2012 Jul.	Paliya et al. (2013b)
	4.0 ± 0.4	2.60 ± 0.06	305	2008 Aug.–2012 Nov.	D’Ammando et al. (2013a)
J2007 – 4434	1.2 ± 0.3	2.3 ± 0.1	44	2008 Aug.–2011 Feb.	Foschini (2011a)
	1.7 ± 0.3	2.6 ± 0.1	68	2008 Aug.–2012 Jul.	Paliya et al. (2013b)
	1.4 ± 0.3	2.6 ± 0.1	49	2008 Aug.–2012 Aug.	D’Ammando et al. (2013d)

Notes. Columns: (1) name of the source; (2) flux in the 0.1–100 GeV energy band [10^{-8} ph cm $^{-2}$ s $^{-1}$]; (3) photon index of the power-law model ($\Gamma = \alpha + 1$); (4) test Statistic (Mattox et al. 1996); (5) time period; (6) reference.

Table 5. X-ray characteristics.

Name	Sat.	Date	Exp.	Γ_1	E_{break}	Γ_2	Stat/val/d.o.f.	$F_{0.3-10 \text{ keV}}$	$L_{0.3-10 \text{ keV}}$
J0100 – 0200	C	2003 Sep. 04 11:46	9.2	2.0(f)			CXA	0.0051 ± 0.0023	0.0090
J0134 – 4258	S	2007 Nov. 29 07:04	2.3	2.4 ± 0.4			$\chi^2/4.3/3$	2.2 ± 0.5	4.6
	S	2008 Mar. 13 01:29	4.0	1.9 ± 0.2			$\chi^2/6.5/7$	2.7 ± 0.5	4.8
	S	2008 Mar. 25 04:27	4.3	2.1 ± 0.2			$\chi^2/1.4/4$	1.5 ± 0.3	2.5
	S	2008 Mar. 27 04:38	5.0	1.5 ± 0.4			$\chi^2/3.8/3$	1.7 ± 0.6	2.7
	S	2008 Mar. 29 01:24	3.3	2.3 ± 0.2			$\chi^2/3.4/6$	2.9 ± 0.4	5.7
	S	2008 Mar. 31 06:29	3.2	1.7 ± 0.2			$\chi^2/8.5/9$	4.6 ± 0.7	7.8
	X	2008 Dec. 11 20:03	32.4	$1.94^{+0.03}_{-0.09}$	$1.4^{+0.1}_{-0.4}$	$2.29^{+0.05}_{-0.08}$	$\chi^2/540.4/541$	2.9 ± 0.8	5.3
	S	2010 Mar. 21 04:38	2.3	2.4 ± 0.3			C/48.1/64	1.3 ± 0.3	2.6
	S	2010 Mar. 21 06:14	1.7	2.1 ± 0.3			C/48.2/61	1.7 ± 0.4	3.1
	S	2010 Nov. 21 17:06	1.4	1.6 ± 0.3			C/60.3/64	2.5 ± 0.9	4.1
	S	2010 Nov. 21 21:55	1.0	1.9 ± 0.3			C/44.3/55	2.7 ± 0.8	4.7
	S	2011 Mar. 20 02:01	1.4	1.7 ± 0.2			C/86.0/107	4.4 ± 0.9	7.4
	S	2011 Mar. 20 03:57	1.3	1.9 ± 0.2			C/87.2/98	4.0 ± 0.9	6.9
	S	2011 Mar. 25 05:32	3.7	1.9 ± 0.2			$\chi^2/7.0/8$	3.5 ± 0.6	6.1
	J0324 + 3410	S	2011 Mar. 26 07:06	2.7	2.0 ± 0.3			$\chi^2/3.8/5$	3.4 ± 0.7
S		2012 Dec. 06 02:38	2.7	2.3 ± 0.3			$\chi^2/6.9/6$	3.6 ± 0.7	7.2
S		2006 Jul. 06 00:53	8.3	2.03 ± 0.06			$\chi^2/99.1/82$	16.1 ± 0.9	1.9
S		2006 Jul. 09 10:51	2.4	2.0 ± 0.1			$\chi^2/38.2/27$	15.6 ± 1.4	1.8
S		2006 Jul. 09 23:42	8.8	2.3 ± 0.1	$1.3^{+0.4}_{-0.2}$	1.8 ± 0.1	$\chi^2/126.3/99$	28.2 ± 9.0	3.3
S		2007 Jul. 20 16:59	6.4	2.04 ± 0.05			$\chi^2/83.9/81$	19.2 ± 0.9	2.3
S		2007 Nov. 04 03:08	1.9	1.9 ± 0.1			$\chi^2/20.9/12$	12.5 ± 1.4	1.4
S		2007 Nov. 11 02:08	2.2	2.0 ± 0.1			$\chi^2/9.9/13$	10.5 ± 1.3	1.2
S		2007 Nov. 25 18:02	2.2	1.9 ± 0.2			$\chi^2/22.1/10$	7.7 ± 1.2	0.87
S		2007 Dec. 01 04:14	2.4	2.0 ± 0.1			$\chi^2/27.1/23$	14.9 ± 1.3	1.8
S		2007 Dec. 06 00:18	6.1	1.92 ± 0.07			$\chi^2/50.0/47$	11.8 ± 0.6	1.3
S		2007 Dec. 15 09:05	2.5	2.1 ± 0.1			$\chi^2/11.0/13$	9.6 ± 1.2	1.2
S		2007 Dec. 23 08:18	2.1	2.3 ± 0.1			$\chi^2/13.9/24$	15.0 ± 1.1	2.0
S		2007 Dec. 28 06:48	1.3	2.2 ± 0.2			$\chi^2/5.4/9$	10.8 ± 1.6	1.4
S		2008 Jan. 04 04:25	2.3	2.0 ± 0.1			$\chi^2/21.0/21$	13.5 ± 1.3	1.6
S		2008 Jan. 14 08:28	2.3	2.7 ± 0.2	$1.3^{+0.4}_{-0.3}$	$1.8^{+0.1}_{-0.2}$	$\chi^2/43.2/32$	25.1 ± 8.6	3.3
S		2008 Nov. 16 17:35	5.9	$2.5^{+0.5}_{-0.2}$	$1.3^{+0.3}_{-0.4}$	$1.7^{+0.2}_{-0.1}$	$\chi^2/38.8/30$	13.1 ± 5.7	1.6
S		2009 Jul. 24 03:07	3.2	1.71 ± 0.08			$\chi^2/53.2/38$	19.2 ± 1.4	2.0
S		2009 Jul. 27 08:13	2.8	2.1 ± 0.1			$\chi^2/55.5/30$	14.4 ± 1.3	1.7
S		2009 Jul. 30 08:29	2.9	1.87 ± 0.07			$\chi^2/47.3/44$	22.5 ± 1.5	2.5
S		2009 Aug. 02 00:46	3.2	2.4 ± 0.2	$1.8^{+1.2}_{-0.3}$	$1.6^{+0.2}_{-0.7}$	$\chi^2/50.0/32$	16.8 ± 13.5	2.0
S		2009 Aug. 05 00:50	3.3	2.03 ± 0.08			$\chi^2/28.0/33$	14.0 ± 0.9	1.7
S		2009 Aug. 08 12:30	2.9	2.2 ± 0.1			$\chi^2/28.6/31$	14.3 ± 1.1	1.8
S		2010 Oct. 28 21:36	2.9	1.8 ± 0.1			$\chi^2/24.1/29$	16.0 ± 1.6	1.8
S		2010 Oct. 29 02:16	2.9	1.9 ± 0.1			$\chi^2/18.4/19$	11.6 ± 1.0	1.3
S		2010 Oct. 30 02:21	3.1	$2.4^{+0.3}_{-0.2}$	1.5 ± 0.4	1.6 ± 0.2	$\chi^2/36.6/29$	16.4 ± 8.0	2.0
S		2010 Oct. 31 13:40	3.1	1.90 ± 0.09			$\chi^2/33.7/33$	15.8 ± 1.1	1.8
S		2010 Nov. 01 07:20	3.3	2.14 ± 0.09			$\chi^2/66.4/34$	15.1 ± 1.2	1.9
S		2010 Nov. 02 04:18	3.4	2.0 ± 0.1			$\chi^2/46.0/29$	12.7 ± 1.0	1.5
S		2010 Nov. 03 09:07	3.3	2.1 ± 0.1			$\chi^2/32.9/31$	12.8 ± 0.9	1.6
S		2010 Nov. 04 20:42	2.9	1.9 ± 0.1			$\chi^2/25.1/25$	13.1 ± 1.1	1.1
S		2010 Nov. 05 20:46	2.8	2.01 ± 0.08			$\chi^2/39.8/39$	19.6 ± 1.4	2.3
S		2010 Nov. 06 20:46	3.0	1.97 ± 0.09			$\chi^2/37.9/31$	15.8 ± 1.4	1.8
S	2010 Nov. 07 20:57	1.4	2.1 ± 0.2			$\chi^2/9.8/5$	9.2 ± 1.7	1.1	
S	2010 Nov. 08 21:00	2.9	2.0 ± 0.1			$\chi^2/26.8/20$	10.9 ± 0.9	1.3	
S	2010 Nov. 09 19:32	2.9	2.1 ± 0.1			$\chi^2/29.7/29$	14.4 ± 1.2	1.8	
S	2010 Nov. 10 21:12	2.4	2.0 ± 0.1			$\chi^2/27.0/23$	16.3 ± 1.4	1.9	
S	2010 Nov. 11 13:00	3.1	2.07 ± 0.07			$\chi^2/51.5/45$	22.0 ± 1.3	2.7	
S	2010 Nov. 12 01:54	2.8	2.1 ± 0.1			$\chi^2/34.4/29$	18.3 ± 1.4	2.2	
S	2010 Nov. 13 13:10	3.0	2.0 ± 0.1			$\chi^2/30.2/29$	22.7 ± 1.9	2.6	
S	2010 Nov. 14 05:16	3.0	1.93 ± 0.07			$\chi^2/64.3/47$	23.3 ± 1.5	2.7	
S	2010 Nov. 15 00:47	1.9	1.9 ± 0.2			$\chi^2/6.2/8$	11.5 ± 2.1	1.3	
S	2010 Nov. 16 13:40	3.1	1.80 ± 0.08			$\chi^2/52.6/38$	22.3 ± 1.7	2.4	

Notes. Columns: (1) name of the source; (2) satellite used: C, *Chandra*; S, *Swift*; X, *XMM-Newton*; R, ROSAT; (3) observing date (start, YYYY/MM/DD HH:MM; or interval of dates, if it is the result of an integration of different snapshots); (4) exposure [ks]; (5) photon index (power-law model, $\Gamma = \alpha + 1$) or low-energy photon index (broken power-law model); (6) break energy [keV]; (7) high-energy photon index (broken power-law model); (8) statistics/value/degrees of freedom, where statistics can be χ^2 , or likelihood (Cash 1979), or CXA, XSS, or RASS, if the detection is extracted from the *Chandra*, *XMM-Newton*, or ROSAT catalogues; (9) observed flux in the 0.3–10 keV [10^{-12} erg cm $^{-2}$ s $^{-1}$]; (10) intrinsic luminosity [10^{44} erg s $^{-1}$].

Table 5. continued.

Name	Sat.	Date	Exp.	Γ_1	E_{break}	Γ_2	Stat/val/d.o.f.	$F_{0.3-10 \text{ keV}}$	$L_{0.3-10 \text{ keV}}$
	S	2010 Nov. 17 11:54	3.0	1.9 ± 0.1			$\chi^2/29.4/24$	20.9 ± 1.9	2.4
	S	2010 Nov. 18 07:14	2.8	1.97 ± 0.09			$\chi^2/49.9/31$	16.3 ± 1.4	1.9
	S	2010 Nov. 19 15:50	3.0	1.9 ± 0.1			$\chi^2/27.1/23$	11.5 ± 1.0	1.3
	S	2010 Nov. 23 04:25	2.6	2.29 ± 0.08			$\chi^2/52.7/45$	22.6 ± 1.4	3.0
	S	2010 Nov. 24 18:57	2.7	2.09 ± 0.08			$\chi^2/33.1/43$	22.7 ± 1.6	2.8
	S	2010 Nov. 25 19:03	2.9	2.4 ± 0.1	$2.0^{+0.7}_{-0.4}$	$1.6^{+0.2}_{-0.3}$	$\chi^2/42.2/43$	25.5 ± 10.3	3.2
	S	2010 Nov. 26 06:18	3.4	2.13 ± 0.07			$\chi^2/59.2/51$	19.6 ± 1.2	2.4
	S	2010 Nov. 27 17:39	3.0	2.4 ± 0.2	$1.4^{+0.4}_{-0.2}$	1.6 ± 0.2	$\chi^2/33.3/30$	20.8 ± 6.9	2.5
	S	2010 Nov. 28 03:17	3.3	2.4 ± 0.1	$2.3^{+0.8}_{-0.6}$	$1.4^{+0.4}_{-0.7}$	$\chi^2/35.1/40$	22.2 ± 13.6	2.7
	S	2010 Nov. 29 03:23	3.3	2.01 ± 0.08			$\chi^2/54.3/46$	19.8 ± 1.4	2.4
	S	2010 Nov. 30 11:46	3.3	2.13 ± 0.08			$\chi^2/35.1/38$	17.9 ± 1.2	2.2
	S	2011 Jul. 06 23:59	2.0	2.07 ± 0.08			$\chi^2/52.2/42$	28.2 ± 1.8	3.4
	S	2011 Aug. 07 05:40	2.0	2.05 ± 0.09			$\chi^2/45.7/32$	22.0 ± 1.7	2.6
	S	2011 Sep. 04 00:07	1.9	2.0 ± 0.1			$\chi^2/17.0/19$	18.3 ± 1.7	2.2
	S	2011 Oct. 03 18:10	2.3	2.0 ± 0.1			$\chi^2/15.0/15$	10.7 ± 1.2	1.3
	S	2011 Oct. 07 01:15	1.9	1.9 ± 0.1			$\chi^2/33.9/29$	24.1 ± 2.1	2.7
	S	2011 Nov. 05 04:50	2.0	2.0 ± 0.1			$\chi^2/11.4/22$	17.4 ± 1.7	2.1
	S	2011 Dec. 04 15:18	0.2	1.9 ± 0.4			C/45.6/52	14.3 ± 1.9	1.6
	S	2011 Dec. 07 01:10	1.5	1.8 ± 0.2			$\chi^2/10.8/7$	9.5 ± 1.8	1.0
	S	2011 Dec. 28 08:48	1.9	2.1 ± 0.1			$\chi^2/17.3/20$	21.1 ± 2.1	2.5
	S	2012 Jan. 02 01:03	1.3	2.0 ± 0.2			$\chi^2/10.3/9$	11.1 ± 1.9	1.3
	S	2012 Jan. 30 15:51	2.0	1.8 ± 0.1			$\chi^2/18.6/15$	12.7 ± 1.6	1.4
	S	2012 Mar. 03 08:07	1.5	1.9 ± 0.2			$\chi^2/6.3/7$	8.2 ± 1.5	0.94
	S	2013 Jan. 13 10:57	3.6	1.9 ± 0.1			$\chi^2/30.8/27$	11.6 ± 1.1	1.3
	S	2013 Jan. 14 07:47	2.8	2.0 ± 0.1			$\chi^2/24.2/23$	12.7 ± 1.3	1.5
	S	2013 Jan. 15 09:18	3.7	2.0 ± 0.1			$\chi^2/31.7/30$	11.9 ± 1.0	1.4
	S	2013 Feb. 15 07:42	4.0	1.8 ± 0.1			$\chi^2/15.6/19$	9.5 ± 1.1	1.0
	S	2013 Feb. 15 18:50	3.8	2.0 ± 0.1			$\chi^2/23.2/20$	9.1 ± 0.9	1.1
	S	2013 Mar. 02 16:04	1.9	2.1 ± 0.1			$\chi^2/9.7/14$	11.4 ± 1.2	1.4
	S	2013 Mar. 05 14:35	2.3	$2.6^{+0.3}_{-0.2}$	$2.2^{+0.6}_{-0.8}$	$1.3^{+0.6}_{-0.7}$	$\chi^2/20.6/18$	18.8 ± 12.5	2.4
J0706 + 3901	C	2003 Jan. 22 22:33	9.6	2.0(f)			CXA	0.0082 ± 0.0025	0.0020
J0713 + 3820	S	2011–2012	2.0	2.0 ± 0.2			$\chi^2/1.0/3$	3.4 ± 0.7	1.6
J0744 + 5149	S	2011 Dec. 16–19	4.7	2.4 ± 0.3			C/75.3/73	0.68 ± 0.19	8.0
J0804 + 3853	S	2010–2012	4.1	3.0 ± 0.3			C/76.3/58	0.68 ± 0.15	1.6
J0814 + 5609	S	2011 Dec. 15–16	3.1	2.0 ± 0.4			C/31.4/36	0.62 ± 0.29	7.1
J0849 + 5108	S	2011–2013	42.1	1.7 ± 0.1			$\chi^2/14.8/14$	0.59 ± 0.09	8.0
J0902 + 0443	S	2012 Jun. 04–06	4.6	2.0(f)			UL 3σ	<0.10	<1.3
J0937 + 3615	S	2012 Dec. 18 04:04	3.4	1.3 ± 0.6			C/13.2/20	0.46 ± 0.34	3.7
J0945 + 1915	S	2012–2013	6.4	1.3 ± 0.4			C/42.7/29	0.38 ± 0.23	8.5
J0948 + 0022	X	2008 Apr. 29 14:42	0.05	1.4 ± 0.3			C/49.2/52	9.3 ± 4.3	111.6
	S	2008 Dec. 05 02:21	4.2	1.8 ± 0.2			$\chi^2/4.0/8$	3.6 ± 0.8	54.0
	S	2009 Mar. 26 06:21	4.8	1.7 ± 0.1			$\chi^2/24.3/23$	7.5 ± 0.9	107.4
	S	2009 Apr. 15 08:16	4.4	1.7 ± 0.1			$\chi^2/11.3/13$	5.2 ± 0.9	73.3
	S	2009 May 05 10:06	4.8	1.6 ± 0.1			$\chi^2/17.6/22$	7.6 ± 1.1	99.6
	S	2009 May 10 12:21	4.9	1.8 ± 0.1			$\chi^2/12.8/12$	4.2 ± 0.8	62.7
	S	2009 May 15 03:04	1.4	1.8 ± 0.3			C/51.2/57	2.1 ± 0.4	31.2
	S	2009 May 25 10:25	5.0	1.7 ± 0.1			$\chi^2/6.9/11$	4.0 ± 0.6	58.1
	S	2009 Jun. 04 03:56	4.5	1.7 ± 0.2			$\chi^2/13.9/10$	4.1 ± 1.0	56.8
	S	2009 Jun. 14 01:04	3.9	1.6 ± 0.1			$\chi^2/11.0/11$	5.3 ± 1.0	70.7
	S	2009 Jun. 23 10:03	7.7	1.6 ± 0.1			$\chi^2/10.5/15$	3.5 ± 0.4	46.0
	S	2009 Jun. 24 19:29	4.7	1.5 ± 0.2			$\chi^2/5.4/7$	3.3 ± 0.9	40.6
	S	2009 Jul. 03 12:41	4.2	1.7 ± 0.3			$\chi^2/4.2/5$	3.0 ± 1.0	41.2
	S	2010 Jul. 03 19:30	1.6	1.4 ± 0.3			$\chi^2/1.1/2$	5.6 ± 2.2	65.0
	S	2011 Apr. 29 04:10	2.0	1.6 ± 0.4			$\chi^2/0.82/2$	4.0 ± 1.7	52.5
	S	2011 May 15 03:53	4.7	1.8 ± 0.2			$\chi^2/15.8/10$	4.0 ± 0.9	58.0
	S	2011 May 28 01:54	3.6	1.6 ± 0.2			$\chi^2/12.5/9$	4.6 ± 1.1	62.1
	X	2011 May 28 11:21	34.9	2.33 ± 0.03	1.21 ± 0.06	1.61 ± 0.03	$\chi^2/943.5/932$	3.9 ± 0.2	78.5
	S	2011 Jun. 04 08:54	2.0	1.7 ± 0.3			$\chi^2/0.95/2$	3.3 ± 1.3	45.7
	S	2011 Jun. 14 03:17	5.2	1.7 ± 0.2			$\chi^2/5.7/9$	3.2 ± 0.6	45.6
	S	2011 Jul. 02 17:30	2.0	1.4 ± 0.2			$\chi^2/5.8/7$	5.0 ± 1.5	59.3
	S	2011 Oct. 09 00:12	0.2	1.4 ± 0.5			C/39.5/33	10.2 ± 8.6	117.4
	S	2011 Oct. 12 21:11	1.9	1.6 ± 0.3			$\chi^2/1.5/4$	10.0 ± 3.2	132.7
	S	2011 Nov. 05 01:57	2.0	1.6 ± 0.2			$\chi^2/3.8/6$	7.1 ± 1.8	95.6
	S	2011 Dec. 08 23:59	1.9	1.5 ± 0.2			C/80.7/95	3.3 ± 1.1	41.7
	S	2012 Jan. 02 11:08	2.0	1.3 ± 0.4			$\chi^2/0.72/1$	3.6 ± 1.1	40.5

Table 5. continued.

Name	Sat.	Date	Exp.	Γ_1	E_{break}	Γ_2	Stat/val/d.o.f.	$F_{0.3-10 \text{ keV}}$	$L_{0.3-10 \text{ keV}}$
	S	2012 Jan. 05 23:59	0.5	1.7 ± 0.5			C/22.1/21	2.8 ± 2.1	39.5
	S	2012 Jan. 31 05:21	0.8	1.7 ± 0.3			C/42.5/49	3.4 ± 1.6	46.4
	S	2012 Feb. 27 14:56	2.1	1.8 ± 0.5			$\chi^2/0.34/1$	2.6 ± 1.6	38.9
	S	2012 Mar. 26 03:38	2.0	1.8 ± 0.2			$\chi^2/6.0/5$	5.3 ± 1.3	78.9
	S	2012 Mar. 30 02:16	2.2	1.8 ± 0.3			$\chi^2/3.8/4$	4.4 ± 1.4	66.0
	S	2012 Apr. 23 05:46	0.9	1.4 ± 0.3			C/66.7/75	5.1 ± 2.0	61.8
	S	2012 Apr. 28 13:43	2.0	1.4 ± 0.3			C/75.0/65	2.1 ± 1.0	25.7
	S	2012 May 21 11:54	0.7	1.7 ± 0.3			C/46.5/64	4.7 ± 1.8	65.4
	S	2012 Jun. 18 12:03	2.1	1.6 ± 0.3			$\chi^2/0.34/2$	3.8 ± 1.7	50.8
	S	2012 Jun. 30 18:57	2.0	1.8 ± 0.4			$\chi^2/2.0/2$	4.8 ± 2.1	70.6
	S	2012 Nov. 05 01:51	2.1	$1.8^{+0.4}_{-0.3}$			$\chi^2/2.6/2$	3.8 ± 2.0	55.4
	S	2012 Dec. 03 04:56	2.0	$1.8^{+0.6}_{-0.5}$			$\chi^2/0.23/1$	2.3 ± 0.8	33.4
	S	2012 Dec. 25 07:15	1.9	1.4 ± 0.4			$\chi^2/0.11/1$	4.9 ± 3.4	58.6
	S	2012 Dec. 30 15:34	0.7	1.8 ± 0.2			$\chi^2/1.6/4$	13.1 ± 3.8	198.5
	S	2013 Jan. 03 10:52	3.0	1.5 ± 0.1			$\chi^2/8.1/10$	6.9 ± 1.2	85.8
	S	2013 Jan. 11 08:07	2.9	1.6 ± 0.2			$\chi^2/6.4/9$	6.6 ± 1.6	85.2
	S	2013 Jan. 17 08:17	3.3	1.7 ± 0.1			$\chi^2/7.1/13$	9.2 ± 1.6	126.4
J0953 + 2836	S	2012–2013	11.0	1.9 ± 0.3			C/51.6/61	0.28 ± 0.10	5.8
J1031 + 4234	S	2012 Jan.–Oct.	9.2	1.6 ± 0.3			C/56.2/41	0.23 ± 0.10	1.1
J1037 + 0036	S	2013 Jan.–Feb.	6.9	2.0 ± 0.3			C/60.3/50	0.43 ± 0.16	7.5
J1038 + 4227	S	2013 Jan.–Feb.	7.4	1.7 ± 0.5			C/31.4/24	0.16 ± 0.09	0.23
J1047 + 4725	S	2012 Feb.–Apr.	6.7	1.6 ± 0.5			C/20.3/17	0.18 ± 0.16	4.6
J1048 + 2222	S	2013 Jan.–Feb.	3.1	2.0(f)			C/10.1/16	0.25 ± 0.11	0.95
J1102 + 2239	S	2012 Jan.–Jul.	25.1	$1.6 \pm 0.2^*$			$\chi^2/0.94/3$	0.40 ± 0.19	1.5
	X	2012 Jun. 11 07:00	5.0	2.0 ± 0.2			$\chi^2/7.8/11$	0.20 ± 0.04	1.7
J1110 + 3653	S	2012 Mar.–Oct.	9.0	1.0 ± 0.5			C/21.2/17	0.16 ± 0.12	1.8
J1138 + 3653	S	2012–2013	7.4	1.5 ± 0.5			C/23.4/15	0.33 ± 0.22	1.3
J1146 + 3236	S	2013 Jan.–Apr.	7.4	2.1 ± 0.3			C/68.7/72	0.44 ± 0.11	4.0
J1159 + 2838	S	2012 Mar.–Oct.	4.3	1.8 ± 0.6			C/17.8/12	0.16 ± 0.11	0.21
J1227 + 3214	S	2012 Feb.–Aug.	4.0	1.3 ± 0.2			$\chi^2/8.8/6$	2.6 ± 0.7	1.2
J1238 + 3942	S	2012 Nov. 10–14	5.0	2.0 ± 0.3			C/43.6/62	0.66 ± 0.23	11.5
J1246 + 0238	S	2012–2013	23.6	1.6 ± 0.3			$\chi^2/1.0/3$	0.23 ± 0.08	0.95
	X	2012 Dec. 14 21:05	10.3	2.0 ± 0.2			$\chi^2/25.5/16$	0.13 ± 0.02	0.76
J1333 + 4141	X	2006 Dec. 19 09:57	0.0055	2.0(f)			XSS	3.0 ± 0.9	4.7
J1346 + 3121	R	1990–1991	0.714	2.0(f)			RASS	<0.11	<0.20
J1348 + 2622	X	2003 Jan. 13 13:25	40.1	3.32 ± 0.04			$\chi^2/314.5/253$	0.379 ± 0.008	58.8
J1358 + 2658	R	1990–1991	0.531	2.0(f)			RASS	<0.16	<0.63
J1421 + 2824	S	2012 Aug. 16–17	6.1	1.6 ± 0.3			C/53.7/44	0.41 ± 0.18	4.1
J1505 + 0326	S	2009–2012	24.2	1.8 ± 0.2			$\chi^2/8.5/7$	0.32 ± 0.07	2.0
	X	2012 Aug. 07 20:12	10.8	2.1 ± 0.1	1.8 ± 0.5	$1.4^{+0.2}_{-0.3}$	$\chi^2/47.4/49$	0.37 ± 0.13	3.0
J1548 + 3511	X	2011 Aug. 08 00:55	13.9	$2.64^{+0.09}_{-0.08}$	$1.8^{+0.3}_{-0.4}$	1.7 ± 0.2	$\chi^2/95.3/94$	0.41 ± 0.11	6.5
	X	2011 Aug. 20 00:04	17.6	2.55 ± 0.08	1.6 ± 0.3	1.7 ± 0.2	$\chi^2/103.7/126$	0.48 ± 0.11	6.7
J1612 + 4219	R	1990–1991	1.0	2.0(f)			RASS	<0.077	<0.13
J1629 + 4007	S	2005 Apr. 20 07:06	4.9	2.4 ± 0.2			$\chi^2/15.9/11$	2.6 ± 0.3	7.2
	S	2005 May 23 20:38	0.6	2.7 ± 0.4			C/26.8/35	1.9 ± 0.6	5.7
	S	2006 Jan. 20 23:59	6.2	2.2 ± 0.1			$\chi^2/17.3/13$	2.3 ± 0.2	5.9
	S	2007 Apr. 22 05:25	1.9	2.6 ± 0.3			$\chi^2/2.1/2$	2.1 ± 0.4	6.1
	S	2007 Apr. 28 04:45	1.9	2.6 ± 0.4			$\chi^2/1.1/1$	1.9 ± 0.4	5.7
	S	2007 May 04 11:17	2.2	2.3 ± 0.3			C/27.8/43	0.79 ± 0.24	2.1
	S	2007 May 12 08:50	2.1	2.4 ± 0.4			$\chi^2/1.1/2$	2.2 ± 0.5	6.0
	S	2007 May 18 09:36	2.2	2.5 ± 0.5			$\chi^2/1.4/1$	1.8 ± 0.5	5.1
	S	2007 May 27 05:39	2.2	1.9 ± 0.3			$\chi^2/0.91/2$	2.7 ± 0.7	6.5
	S	2007 Jun. 26 00:32	3.0	2.4 ± 0.3			$\chi^2/0.71/2$	2.0 ± 0.4	5.6
	S	2008 Jan. 08 01:03	7.6	2.4 ± 0.1			$\chi^2/25.1/14$	2.8 ± 0.3	7.6
	S	2008 Jan. 11 19:09	6.5	2.5 ± 0.2			$\chi^2/4.7/8$	1.8 ± 0.2	5.0
	S	2008 Apr. 05 01:52	2.0	2.4 ± 0.4			$\chi^2/1.8/1$	1.6 ± 0.5	4.4
	S	2008 May 03 10:50	1.0	2.4 ± 0.3			C/48.2/54	2.5 ± 0.6	6.9
	S	2012 Jul. 26 08:22	0.9	1.7 ± 0.4			C/35.1/41	3.0 ± 1.3	6.8
	S	2012 Oct. 13 02:38	0.6	2.5 ± 0.6			C/19.1/26	2.5 ± 1.0	7.1
	S	2012 Oct. 29 16:14	1.4	3.6 ± 0.4			C/19.8/40	1.7 ± 0.5	6.5
	S	2012 Oct. 30 11:45	1.1	2.8 ± 0.4			C/25.9/38	1.6 ± 0.5	4.8
J1633 + 4718	X	2011 Jul. 09 05:50	17.0	1.58 ± 0.05	$1.8^{+0.5}_{-0.3}$	$1.94^{+0.10}_{-0.08}$	$\chi^2/522.0/368$	1.9 ± 0.5	0.79
	X	2011 Sep. 12 22:24	19.0	1.59 ± 0.03	$2.2^{+0.4}_{-0.3}$	$1.96^{+0.10}_{-0.08}$	$\chi^2/582.0/427$	2.0 ± 0.4	0.83
	X	2012 Jan. 14 15:56	12.8	1.40 ± 0.04	$2.2^{+0.4}_{-0.3}$	$1.96^{+0.11}_{-0.09}$	$\chi^2/483.4/370$	2.5 ± 0.5	0.98
	X	2012 Mar. 14 10:20	8.7	1.60 ± 0.03			$\chi^2/316.0/211$	1.96 ± 0.07	0.76

Table 5. continued.

Name	Sat.	Date	Exp.	Γ_1	E_{break}	Γ_2	Stat/val/d.o.f.	$F_{0.3-10 \text{ keV}}$	$L_{0.3-10 \text{ keV}}$
	S	2012 May–Jun.	1.2	2.0 ± 0.4			C/27.5/41	1.6 ± 0.6	0.61
J1634 + 4809	S	2011–2012	6.6	1.8 ± 0.4			C/38.7/35	0.26 ± 0.12	2.4
J1644 + 2619	S	2011 Dec. 26 09:41	1.3	2.2 ± 0.3			C/73.9/63	2.1 ± 0.6	1.5
J1709 + 2348	R	1990–1991	8.0	2.0(f)			RASS	0.65 ± 0.06	1.3
J2007 – 4434	X	2004 Apr. 11 18:52	20.7	2.1 ± 0.2	$0.65^{+0.10}_{-0.08}$	1.53 ± 0.03	$\chi^2/351.1/338$	1.4 ± 0.4	3.9
	S	2011–2013	58.6	1.7 ± 0.1			$\chi^2/13.4/23$	0.60 ± 0.06	1.0
	X	2012 May 01 05:32	19.0	1.69 ± 0.05			$\chi^2/166.0/146$	0.47 ± 0.02	1.2
	X	2012 Oct. 18 21:24	23.2	1.67 ± 0.03			$\chi^2/235.2/206$	0.69 ± 0.02	1.5
J2021 – 2235	X	2005 Apr. 06 14:31	0.011	2.0(f)			XSS	2.1 ± 0.6	2.5

Table 6. *Swift*/UVOT observed average magnitudes (extracted from all the data integrated).

Name	v	Exp.	b	Exp.	u	Exp.	$uw1$	Exp.	$uw2$	Exp.	$uw2$	Exp.
J0134 – 4258	16.21 ± 0.03	2.5	16.35 ± 0.03	2.6	15.17 ± 0.03	5.1	14.73 ± 0.04	5.3	14.59 ± 0.04	7.0	14.69 ± 0.04	16.0
J0324 + 3410	15.70 ± 0.02	41.0	16.24 ± 0.03	15.5	15.34 ± 0.03	17.0	15.61 ± 0.04	31.8	15.91 ± 0.04	40.1	15.80 ± 0.04	91.7
J0713 + 3820	14.94 ± 0.04	0.08	15.38 ± 0.04	0.12	14.41 ± 0.04	0.12	14.70 ± 0.04	1.1	14.92 ± 0.06	0.11	15.12 ± 0.04	0.47
J0744 + 5149	18.90 ± 0.03	0.17	18.57 ± 0.11	0.17	17.63 ± 0.04	1.2	17.74 ± 0.06	0.76	17.88 ± 0.06	1.6	17.87 ± 0.07	0.67
J0804 + 3853					17.25 ± 0.04	1.7	17.76 ± 0.08	0.64	18.11 ± 0.08	0.94	18.21 ± 0.08	0.79
J0814 + 5609	18.04 ± 0.17	0.18	18.40 ± 0.09	0.18	17.47 ± 0.05	0.52	17.21 ± 0.07	0.36	17.22 ± 0.06	1.1	17.14 ± 0.05	0.72
J0849 + 5108	17.32 ± 0.04	2.4	17.89 ± 0.03	2.2	17.60 ± 0.04	2.2	17.60 ± 0.04	16.8	17.60 ± 0.04	12.9	17.63 ± 0.04	12.6
J0902 + 0443	19.31 ± 0.45	0.27	19.36 ± 0.22	0.27	18.54 ± 0.10	0.58	18.45 ± 0.08	1.5	18.36 ± 0.10	0.82	18.56 ± 0.08	1.1
J0937 + 3615	17.42 ± 0.13	0.01	17.97 ± 0.09	0.01	17.18 ± 0.08	0.01	17.49 ± 0.04	2.3	17.60 ± 0.10	0.31	17.83 ± 0.08	0.40
J0945 + 1915	16.71 ± 0.06	0.24	17.01 ± 0.04	0.24	15.93 ± 0.03	1.5	15.79 ± 0.04	1.5	15.69 ± 0.04	1.9	15.85 ± 0.04	1.6
J0948 + 0022	17.72 ± 0.03	6.6	18.10 ± 0.03	6.7	17.38 ± 0.03	40.1	17.11 ± 0.04	14.2	17.20 ± 0.04	17.4	17.21 ± 0.04	26.9
J0953 + 2836	19.44 ± 0.34	0.37	19.33 ± 0.13	0.37	18.31 ± 0.04	6.4	18.20 ± 0.08	0.74	17.93 ± 0.06	1.4	18.01 ± 0.05	1.5
J1031 + 4234	19.39 ± 0.26	0.72	19.82 ± 0.16	0.72	19.27 ± 0.13	0.72	19.65 ± 0.13	1.6	19.62 ± 0.12	2.1	20.02 ± 0.11	3.0
J1037 + 0036	>19.3	0.31	19.97 ± 0.24	0.31	19.46 ± 0.21	0.31	19.93 ± 0.16	1.4	20.00 ± 0.24	0.83	20.90 ± 0.17	3.6
J1038 + 4227	17.73 ± 0.07	0.53	18.27 ± 0.05	0.53	17.92 ± 0.05	0.67	18.74 ± 0.07	1.6	19.26 ± 0.10	1.8	19.83 ± 0.09	2.7
J1047 + 4725	19.05 ± 0.21	0.35	19.35 ± 0.12	0.35	18.48 ± 0.06	1.3	18.11 ± 0.06	1.6	17.68 ± 0.06	1.5	18.00 ± 0.05	1.9
J1048 + 2222	18.86 ± 0.30	0.17	18.78 ± 0.13	0.17	17.83 ± 0.09	0.17	17.66 ± 0.05	1.3	17.42 ± 0.08	0.52	17.87 ± 0.07	0.69
J1102 + 2239	19.37 ± 0.26	0.76	19.85 ± 0.17	0.76	19.66 ± 0.07	5.0	20.47 ± 0.15	5.3	20.86 ± 0.17	5.5	20.75 ± 0.11	7.4
J1110 + 3653	>19.6	0.37	20.51 ± 0.39	0.37	20.10 ± 0.26	0.84	19.65 ± 0.10	3.0	19.67 ± 0.10	3.4	19.77 ± 0.11	2.1
J1138 + 3653	19.25 ± 0.20	0.63	20.08 ± 0.17	0.63	19.64 ± 0.16	0.63	20.37 ± 0.17	2.3	21.29 ± 0.38	2.0	21.60 ± 0.32	2.7
J1146 + 3236	18.73 ± 0.15	0.49	19.11 ± 0.49	0.49	18.33 ± 0.05	1.1	18.34 ± 0.06	1.3	18.27 ± 0.07	1.3	18.33 ± 0.05	2.5
J1159 + 2838	18.13 ± 0.16	0.24	18.72 ± 0.12	0.24	17.82 ± 0.05	0.93	18.23 ± 0.07	1.0	18.44 ± 0.12	0.49	18.66 ± 0.07	1.4
J1227 + 3214	18.22 ± 0.37	0.07	19.01 ± 0.36	0.07	19.07 ± 0.08	2.3	19.14 ± 0.20	0.33	19.88 ± 0.30	0.45	19.48 ± 0.14	0.82
J1238 + 3942	>18.9	0.21	>19.8	0.21	18.99 ± 0.24	0.21	18.65 ± 0.13	0.42	18.54 ± 0.16	0.34	18.66 ± 0.05	3.5
J1246 + 0238	18.22 ± 0.09	0.71	18.46 ± 0.06	0.71	17.61 ± 0.03	4.3	17.53 ± 0.04	7.6	17.56 ± 0.05	2.8	17.64 ± 0.04	7.5
J1421 + 2824	17.63 ± 0.13	0.13	17.88 ± 0.08	0.17	16.87 ± 0.06	0.18	16.69 ± 0.04	4.7	16.70 ± 0.07	0.28	16.64 ± 0.05	0.51
J1505 + 0326	19.02 ± 0.13	1.0	19.65 ± 0.11	1.0	18.68 ± 0.04	4.4	18.42 ± 0.04	6.3	18.67 ± 0.05	6.6	18.44 ± 0.05	4.2
J1629 + 4007	18.18 ± 0.05	3.0	18.37 ± 0.04	2.6	17.32 ± 0.04	9.1	17.04 ± 0.03	7.9	16.82 ± 0.04	11.7	16.84 ± 0.04	10.8
J1633 + 4718					16.47 ± 0.05	0.27	16.83 ± 0.06	0.30	16.89 ± 0.06	0.59		
J1634 + 4809	>19.1	0.14	20.00 ± 0.35	0.14	19.43 ± 0.30	0.14	19.82 ± 0.16	1.1	19.96 ± 0.11	4.1	20.15 ± 0.19	0.90
J1644 + 2619	17.88 ± 0.26	0.08	18.22 ± 0.14	0.08	16.83 ± 0.08	0.08	16.98 ± 0.08	0.17	16.90 ± 0.09	0.20	16.95 ± 0.05	0.70
J2007 – 4434	18.77 ± 0.06	6.1	19.49 ± 0.05	6.5	18.80 ± 0.03	38.2	19.40 ± 0.04	36.1	19.92 ± 0.06	21.2	20.10 ± 0.06	14.2

Notes. Column Exp. indicated the resulting exposure [ks].

Table 7. Characteristics at radio frequencies (VLBI).

Frequency	Date	Core size	Core flux	Total flux C	Total flux M	RMS C	RMS M	Axes ratio	Pos angle	Direction	Width	EVPA
J0324 + 3410												
15 GHz	2010 Oct. 15	0.187	0.222	0.300	0.307	0.153	0.177	0.3	-51.5	-49.9	0.070	35.0
	2011 Mar. 05	0.056	0.337	0.400	0.407	0.155	0.198	1.0		-61.6	0.056	26.0
	2011 May 26	0.029	0.237	0.336	0.344	0.171	0.185	1.0		-46.5	0.029	31.0
	2011 Jul. 15		0.169	0.286	0.294	0.156	0.183	1.0		-50.8		31.0
	2011 Dec. 29		0.232	0.373	0.379	0.182	0.225	1.0		-53.4		27.0
	2012 Jul. 12	0.061	0.153	0.229	0.237	0.166	0.197	1.0		-50.9	0.061	33.0
	2012 Dec. 23		0.146	0.258	0.265	0.171	0.185	1.0		-57.8		25.0
	2013 Jul. 08	0.027	0.131	0.213	0.219	0.161	0.174	1.0		-53.3	0.027	33.0
J0814 + 5609												
2 GHz	2010 Mar. 23	1.338	0.024	0.029	0.031	0.227	0.222	0.0	78.5	-96.3	1.274	
8 GHz	2010 Mar. 23	0.282	0.034	0.032	0.036	0.199	0.153	0.2	-84.3	-89.5	0.280	
J0849 + 5108												
15 GHz	2011 May 26	0.009	0.183	0.242	0.244	0.168	0.206	1.0		-77.9	0.009	163.0
	2011 Jul. 15	0.060	0.165	0.211	0.215	0.168	0.180	1.0		-82.5	0.060	105.0
	2012 Jan. 02	0.064	0.186	0.229	0.233	0.155	0.182	1.0		-79.4	0.064	174.0
	2012 Jun. 25	0.061	0.265	0.305	0.308	0.105	0.131	1.0		-80.9	0.061	169.0
	2012 Nov. 02	0.034	0.281	0.310	0.314	0.132	0.151	1.0		-65.8	0.034	159.0
	2013 Jan. 21	0.050	0.218	0.247	0.252	0.125	0.136	1.0		-78.9	0.050	168.0
	2013 Jul. 08	0.045	0.333	0.357	0.361	0.153	0.165	1.0		-76.7	0.045	24.0
J0902 + 0443												
5 GHz	2013 Apr. 08	1.040	0.071	0.076	0.109	4.210	1.014	0.3	47.7	-110.5	0.575	
8 GHz	2013 Apr. 08	0.780	0.069	0.051	0.095	4.269	0.383	0.5	18.0	-122.1	0.457	
J0948 + 0022												
15 GHz	2009 May 28	0.025	0.420	0.434	0.436	0.155	0.178	1.0		-157.6	0.025	131.0
	2009 Jul. 23	0.022	0.326	0.338	0.340	0.174	0.175	1.0		-155.2	0.022	146.0
	2009 Dec. 10		0.381	0.383	0.387	0.141	0.155	1.0		-147.2		
	2010 Sep. 17	0.027	0.531	0.531	0.537	0.158	0.188	1.0		-152.3	0.027	51.0
	2010 Nov. 04	0.045	0.487	0.494	0.496	0.168	0.251	1.0		-153.5	0.045	49.0
	2010 Nov. 29	0.029	0.477	0.478	0.483	0.147	0.164	1.0		-151.2	0.029	49.0
	2011 Feb. 20	0.014	0.602	0.619	0.623	0.155	0.207	1.0		-179.0	0.014	19.0
	2011 May 26	0.029	0.652	0.664	0.665	0.181	0.198	1.0		-171.3	0.029	49.0
	2011 Jun. 24	0.023	0.655	0.669	0.671	0.151	0.190	1.0		-158.3	0.023	51.0
	2011 Sep. 12	0.032	0.447	0.455	0.458	0.182	0.195	1.0		-169.9	0.032	40.0
	2011 Dec. 12	0.033	0.366	0.380	0.382	0.181	0.215	1.0		-162.2	0.033	77.0
	2012 Jul. 12	0.056	0.321	0.328	0.330	0.200	0.219	1.0		-171.8	0.056	98.0
	2012 Nov. 11	0.026	0.334	0.338	0.340	0.145	0.164	1.0		-170.1	0.026	82.0
	2012 Dec. 10	0.016	0.379	0.381	0.388	0.139	0.157	1.0		-162.6	0.016	101.0
	2013 Jan. 21	0.033	0.480	0.485	0.487	0.128	0.137	1.0		-172.7	0.033	125.0
	2013 May 05	0.041	0.862	0.866	0.867	0.164	0.194	1.0		-159.4	0.041	102.0
	2013 Jul. 30	0.053	0.420	0.421	0.423	0.188	0.223	1.0		-171.3	0.053	92.0

Notes. Columns: (1) frequency; (2) date; (3) core size [mas]; (4) core flux density [Jy]; (5) total flux density from clean [Jy]; (6) total flux density from model fitting [Jy]; (7) rms after clean [mJy]; (8) rms after model fitting [mJy]; (9) core axial ratio; (10) core position angle; (11) jet direction [deg]; (12) jet width [mas]; (13) MOJAVE EVPA [deg] (Lister et al. 2009, 2013).

Table 7. continued.

Frequency	Date	Core size	Core flux	Total flux C	Total flux M	RMS C	RMS M	Axes ratio	Pos angle	Direction	Width	EVPA
2 GHz	1995 Jul. 15	0.369	0.860	0.855	0.901	2.266	1.836	1.0		85.9	0.369	
	1997 Jan. 10		0.574	0.571	0.587	1.392	1.296	1.0		82.2		
	2000 May 22	0.930	0.496	0.548	0.567	0.931	0.698	0.0	-37.0	104.7	0.352	
	2002 Sep. 25	0.842	0.560	0.579	0.589	1.270	1.101	0.4	-56.0	94.2	0.703	
	2006 Jul. 11	0.851	0.513	0.541	0.557	2.196	1.926	0.2	-25.4	111.9	0.224	
	2006 Sep. 13	0.634	0.389	0.444	0.457	0.975	0.755	0.0	-37.9	120.4	0.083	
	2008 Dec. 17	0.823	0.370	0.376	0.391	1.477	1.209	0.0	-44.1	115.8	0.258	
	2010 Oct. 13	0.591	0.564	0.581	0.598	1.018	0.791	0.7	-46.3	97.6	0.490	
	2012 Oct. 03	0.729	0.538	0.532	0.560	1.572	0.778	0.4	-16.0	102.2	0.290	
	1995 Jul. 15	0.210	0.987	1.002	1.048	2.441	1.279	0.7	-42.3	119.9	0.147	
	1997 Jan. 10	0.014	0.508	0.530	0.558	1.364	0.927	0.0	73.5	108.6	0.014	
	2000 May 22	0.300	0.625	0.648	0.673	1.262	0.626	0.5	-35.0	129.0	0.150	
	2002 Sep. 25	0.192	0.753	0.760	0.784	1.470	0.994	0.0	-47.5	107.1	0.097	
	2006 Jul. 11	0.417	0.580	0.542	0.580	4.709	4.707	0.6	2.6	0.0	0.416	
	2006 Sep. 13	0.235	0.478	0.480	0.495	0.704	0.479	0.5	-55.7	121.1	0.145	
	2008 Dec. 17	0.194	0.502	0.496	0.526	2.361	1.455	0.0	86.3	123.2	0.169	
	2010 Oct. 13	0.321	0.554	0.546	0.568	1.041	1.051	1.0		143.8	0.321	
	2012 Oct. 03	0.192	0.561	0.596	0.616	1.316	0.652	0.0	-39.8	106.4	0.076	
	2013 May 18	0.409	0.429	0.502	0.508	0.427	0.457	0.5	-38.8	139.3	0.232	
	15 GHz	2010 Oct. 15	0.154	0.372	0.502	0.504	0.166	0.201	0.4	-58.1	119.4	0.090
2010 Oct. 25		0.178	0.366	0.485	0.487	0.143	0.188	0.5	-49.8	118.3	0.100	18.0
2011 May 26		0.159	0.348	0.427	0.429	0.187	0.233	0.0	-61.0	118.3	0.086	179.0
2012 Jan. 02		0.122	0.399	0.443	0.447	0.152	0.164	0.3	-64.1	121.2	0.075	50.0
2012 Sep. 27		0.203	0.577	0.584	0.590	0.151	0.318	0.0	-63.7	112.0	0.135	81.0
2013 May 05		0.111	0.401	0.498	0.500	0.160	0.204	0.6	-76.7	120.1	0.094	22.0
2002 Dec. 26			0.456	0.559	0.593	0.838	1.972	1.0		125.2		
2003 May 22		0.124	0.421	0.437	0.452	0.849	2.162	0.0	79.5	112.7	0.122	
2002 Dec. 26		0.042	0.397	0.423	0.447	1.074	1.826	0.0	85.1	107.3	0.041	
8 GHz		2002 Jan. 31	0.425	0.110	0.072	0.110	10.647	9.018	1.0		0.0	0.425

Table 8. Spectral indices ($S_\nu \propto \nu^{-\alpha}$) at different frequencies.

Name	α_r	$\alpha_{5-15\text{GHz}}$	$\alpha_{15-37\text{GHz}}$	$\alpha_{37-142\text{GHz}}$	α_{IR}	α_o	α_{uv}	α_x	α_γ
J0100 – 0200					1.37 ± 0.04	1.18 ± 0.06			
J0134 – 4258*	0.10 ± 0.04				1.14 ± 0.03	0.49 ± 0.02	0.36 ± 0.32	0.98 ± 0.05	
J0324 + 3410*	0.26 ± 0.03	0.24 ± 0.01	-0.57 ± 0.06	0.10 ± 0.06	0.96 ± 0.02	-0.59 ± 0.07	-0.41 ± 0.32	0.83 ± 0.02	1.87 ± 0.09
J0706 + 3901					1.44 ± 0.04	-1.01 ± 0.15			
J0713 + 3820*	0.58 ± 0.12				1.06 ± 0.02	-0.18 ± 0.02	1.43 ± 0.31	1.0 ± 0.2	
J0744 + 5149					1.30 ± 0.05	-0.35 ± 0.17	0.51 ± 0.51	1.4 ± 0.3	
J0804 + 3853					1.21 ± 0.03	-1.75 ± 0.01	1.71 ± 0.65	2.0 ± 0.3	
J0814 + 5609*	0.38 ± 0.01				0.92 ± 0.08	0.40 ± 0.02	-0.22 ± 0.47	1.0 ± 0.4	
J0849 + 5108*	0.21 ± 0.01	-0.25 ± 0.01	-0.64 ± 0.05		0.70 ± 0.09	-1.49 ± 0.02	0.42 ± 0.31	0.7 ± 0.1	1.1 ± 0.1
J0902 + 0443	0.07 ± 0.01				1.45 ± 0.06	-1.20 ± 0.05	0.71 ± 0.65		
J0937 + 3615					1.36 ± 0.03	-1.78 ± 0.02	1.89 ± 0.47	0.3 ± 0.6	
J0945 + 1915					1.19 ± 0.03	-0.94 ± 0.07	0.66 ± 0.31	0.3 ± 0.4	
J0948 + 0022	-0.28 ± 0.01	-0.71 ± 0.01	-0.33 ± 0.01	0.10 ± 0.03	1.33 ± 0.06	-0.39 ± 0.03	0.29 ± 0.32	0.56 ± 0.05	1.7 ± 0.1
J0953 + 2836					0.67 ± 0.18	-0.62 ± 0.01	-0.16 ± 0.53	0.9 ± 0.3	
J1031 + 4234*	-0.40 ± 0.06				1.47 ± 0.69	-1.43 ± 0.04	2.06 ± 0.99	0.6 ± 0.3	
J1037 + 0036					1.00 ± 0.25	-0.89 ± 0.04	3.73 ± 1.41	1.0 ± 0.3	
J1038 + 4227					1.36 ± 0.11	-1.34 ± 0.01	4.77 ± 0.65	0.7 ± 0.5	
J1047 + 4725*	0.33 ± 0.01				1.52 ± 0.13	-0.89 ± 0.05	0.14 ± 0.44	0.6 ± 0.5	
J1048 + 2222					1.20 ± 0.07	-0.70 ± 0.02	1.33 ± 0.48		
J1102 + 2239					1.35 ± 0.04	-1.96 ± 0.04	1.66 ± 1.08	0.8 ± 0.2	2.1 ± 0.2
J1110 + 3653					0.65 ± 0.48	0.20 ± 0.17	0.95 ± 0.87	0.0 ± 0.5	
J1138 + 3653*	0.50 ± 0.09				1.05 ± 0.13	-1.80 ± 0.07	5.27 ± 2.01	0.5 ± 0.5	
J1146 + 3236*	0.38 ± 0.09				0.79 ± 0.23	-0.76 ± 0.05	0.51 ± 0.44	1.1 ± 0.3	
J1159 + 2838					1.76 ± 0.03	-0.04 ± 0.19	2.17 ± 0.56	0.8 ± 0.6	
J1227 + 3214	-1.04 ± 0.07				1.32 ± 0.02	-3.76 ± 0.03	1.87 ± 1.49	0.3 ± 0.2	
J1238 + 3942					0.77 ± 0.19	-0.62 ± 0.28	0.59 ± 0.76	1.0 ± 0.3	
J1246 + 0238	0.55 ± 0.06				1.34 ± 0.11	-0.04 ± 0.04	0.87 ± 0.32	0.8 ± 0.3	2.1 ± 0.3
J1333 + 4141					1.63 ± 0.03	-0.48 ± 0.04			
J1346 + 3121					1.26 ± 0.07	-0.60 ± 0.05			
J1348 + 2622					1.10 ± 0.21	0.00 ± 0.16		2.32 ± 0.04	
J1358 + 2658					1.31 ± 0.04	-0.97 ± 0.07			
J1421 + 2824	-0.21 ± 0.01				1.26 ± 0.06	-0.19 ± 0.03	0.37 ± 0.36	0.6 ± 0.3	
J1505 + 0326	-0.31 ± 0.01	0.05 ± 0.01	0.39 ± 0.07		1.51 ± 0.06	-1.31 ± 0.05	0.21 ± 0.35	0.6 ± 0.2	1.65 ± 0.06
J1548 + 3511*	0.26 ± 0.01				1.31 ± 0.05	-0.06 ± 0.04		0.7 ± 0.2	
J1612 + 4219					1.88 ± 0.03	-0.54 ± 0.07			
J1629 + 4007	-0.68 ± 0.02				1.23 ± 0.04	0.13 ± 0.04	-0.18 ± 0.32	1.26 ± 0.07	
J1633 + 4718	0.42 ± 0.01				1.57 ± 0.02	-0.92 ± 0.01		0.9 ± 0.2	
J1634 + 4809					1.42 ± 0.08	-0.78 ± 0.05	1.80 ± 1.50	0.8 ± 0.4	
J1644 + 2619	-0.06 ± 0.01				1.14 ± 0.07	-1.33 ± 0.01	-0.17 ± 0.53	1.2 ± 0.3	
J1709 + 2348					1.23 ± 0.08	-1.24 ± 0.08			
J2007 – 4434*	0.41 ± 0.01				1.23 ± 0.06	-1.49 ± 0.03	3.03 ± 0.40	0.65 ± 0.05	1.5 ± 0.2
J2021 – 2235*	0.50 ± 0.07				1.61 ± 0.03				

Notes. The index α_r was calculated by using flux densities at 1.4 and 5 GHz when possible; otherwise, we used the available frequencies (mostly 325 MHz and 1.4 GHz, but also 150, 153, 352, and 843 MHz, 8.35 and 37 GHz). An asterisk on the name indicated the detection at frequencies <1.4 GHz.

Table 9. Shortest variability at optical-to- γ ray frequencies.

Name	γ rays	X-rays	<i>uw</i> 2	<i>u</i> <i>m</i> 2	<i>u</i> <i>w</i> 1	<i>u</i>	<i>b</i>	<i>v</i>
J0134 – 4258		-0.071 ± 0.025 (3.8)						
J0324 + 3410	0.10 ± 0.03^a (3.8)	-0.079 ± 0.012 (4.4)	0.09 ± 0.03 (10.7)	0.10 ± 0.04 (7.6)	<0.43 (4.1)	7 ± 3 (6.1)	<0.7 (4.0)	<0.28 (3.6)
J0849 + 5108	12 ± 8^b (4.7)	<18 (3.5)			<0.27 (3.6)			
J0948 + 0022	$<0.8^b$ (5.4)	<0.21 (5.2)	0.12 ± 0.07 (6.3)	0.09 ± 0.05 (6.1)	0.08 ± 0.03 (7.3)	0.07 ± 0.03 (7.9)	0.06 ± 0.03 (7.8)	0.05 ± 0.03 (5.7)
J0953 + 2836					<18 (3.1)			
J1031 + 4234						<0.15 (3.0)		
J1038 + 4227							<12 (3.4)	
J1047 + 4725			<0.83 (3.4)					
J1102 + 2239						<0.24 (3.1)		
J1238 + 3942			<0.58 (3.4)					
J1505 + 0326	1.3 ± 0.5^b (6.6)							<0.04 (3.8)
J1629 + 4007		-0.12 ± 0.02 (3.5)	<0.49 (3.7)			<0.17 (4.1)	<0.23 (3.2)	<0.09 (3.6)
J2007 – 4434	6 ± 2^b (12)	<0.19 (3.3)	<0.12 (3.0)			<0.30 (3.1)		<0.14 (3.1)

Notes. For each source, it is indicated the τ [days] and, in the second row between parentheses, the significance of the flux change [σ].

References. ^(a) Paliya et al. (2014); ^(b) Foschini (2011a).

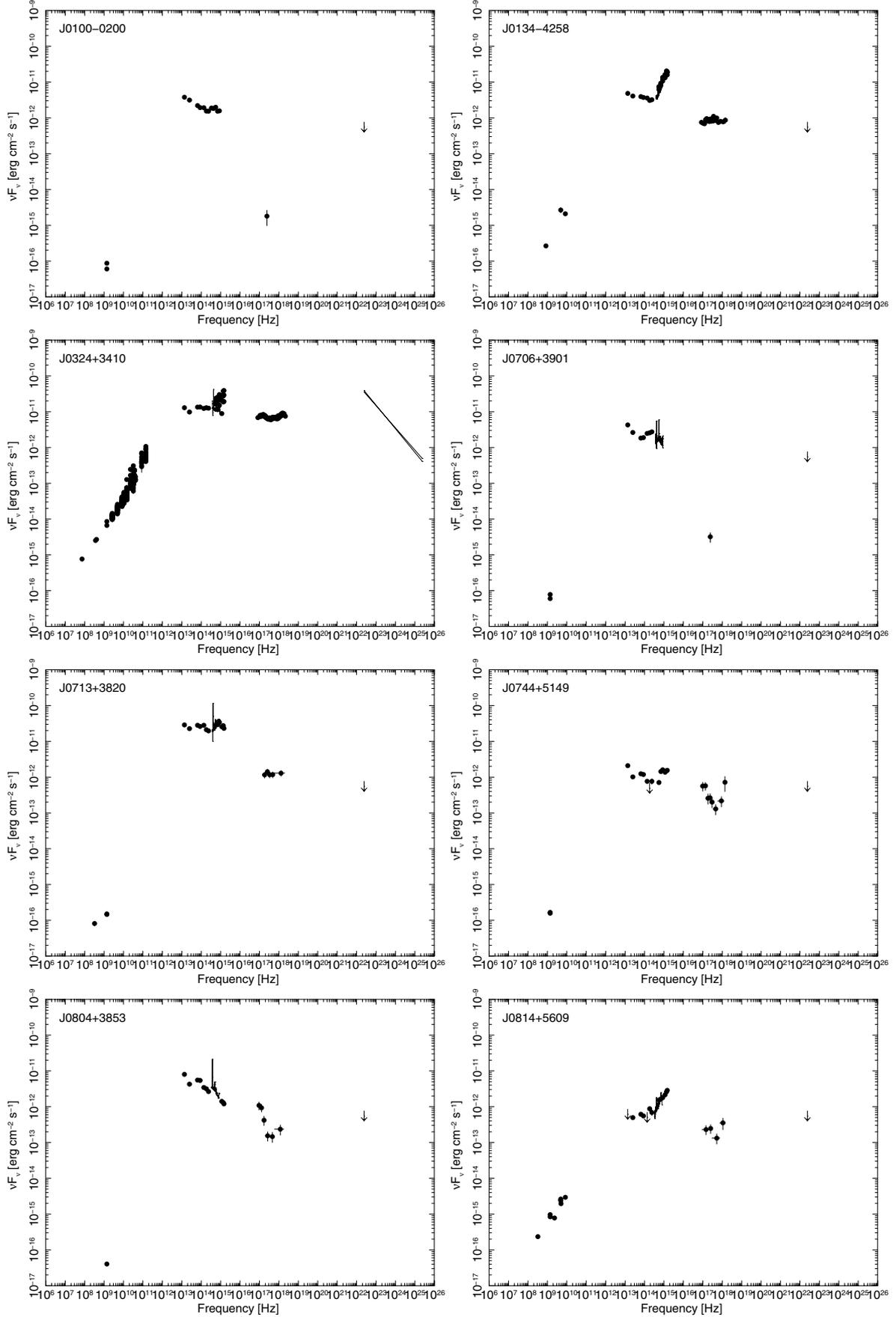


Fig. 8. Spectral energy distributions of the sources in the present sample. Data are corrected for the Galactic absorption. Points refer to detections; arrows are upper limits; the continuous lines are the optical spectra.

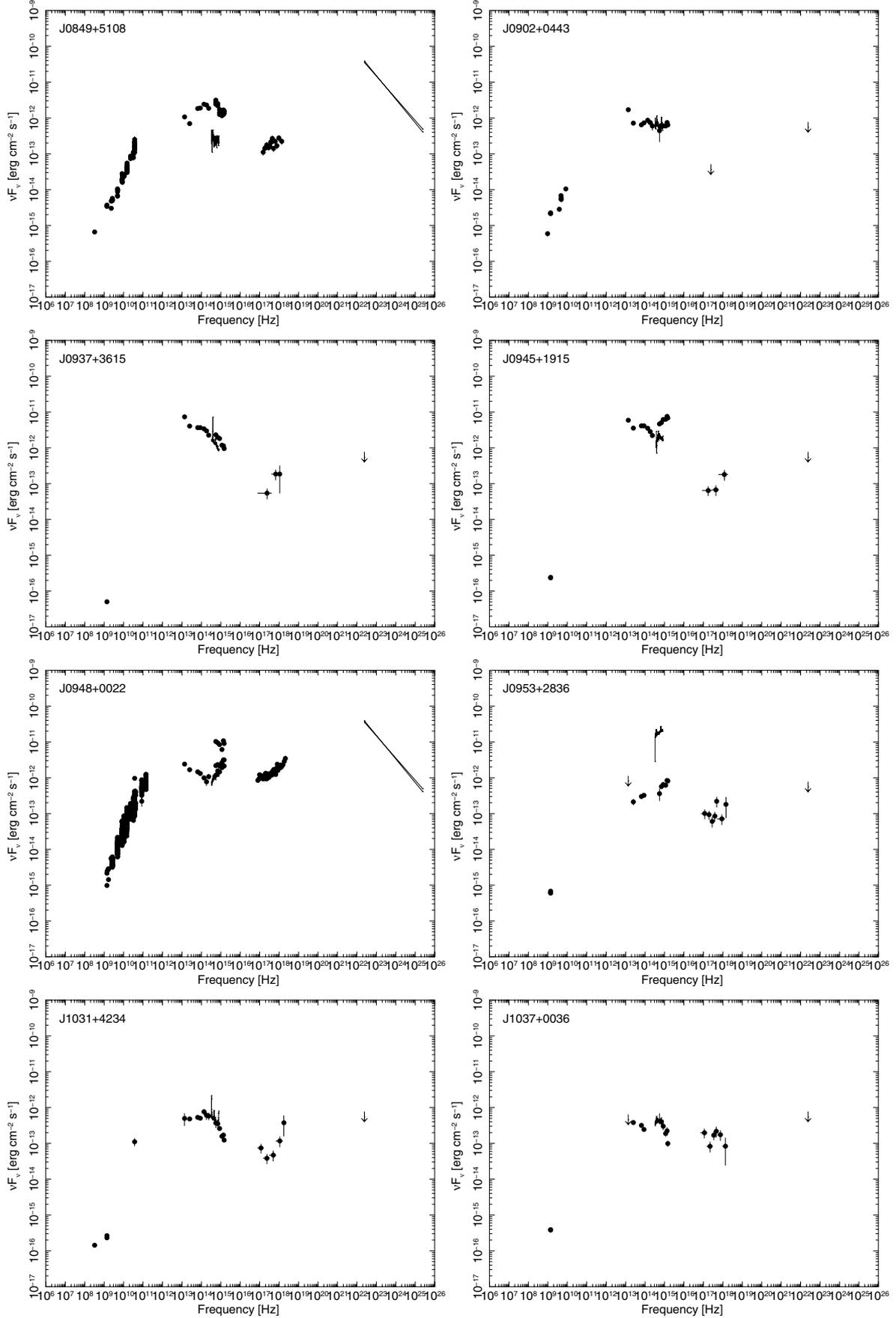


Fig. 9. Spectral energy distributions of the sources in the present sample. Data are corrected for the Galactic absorption. Points refer to detections; arrows are upper limits; the continuous lines are the optical spectra.

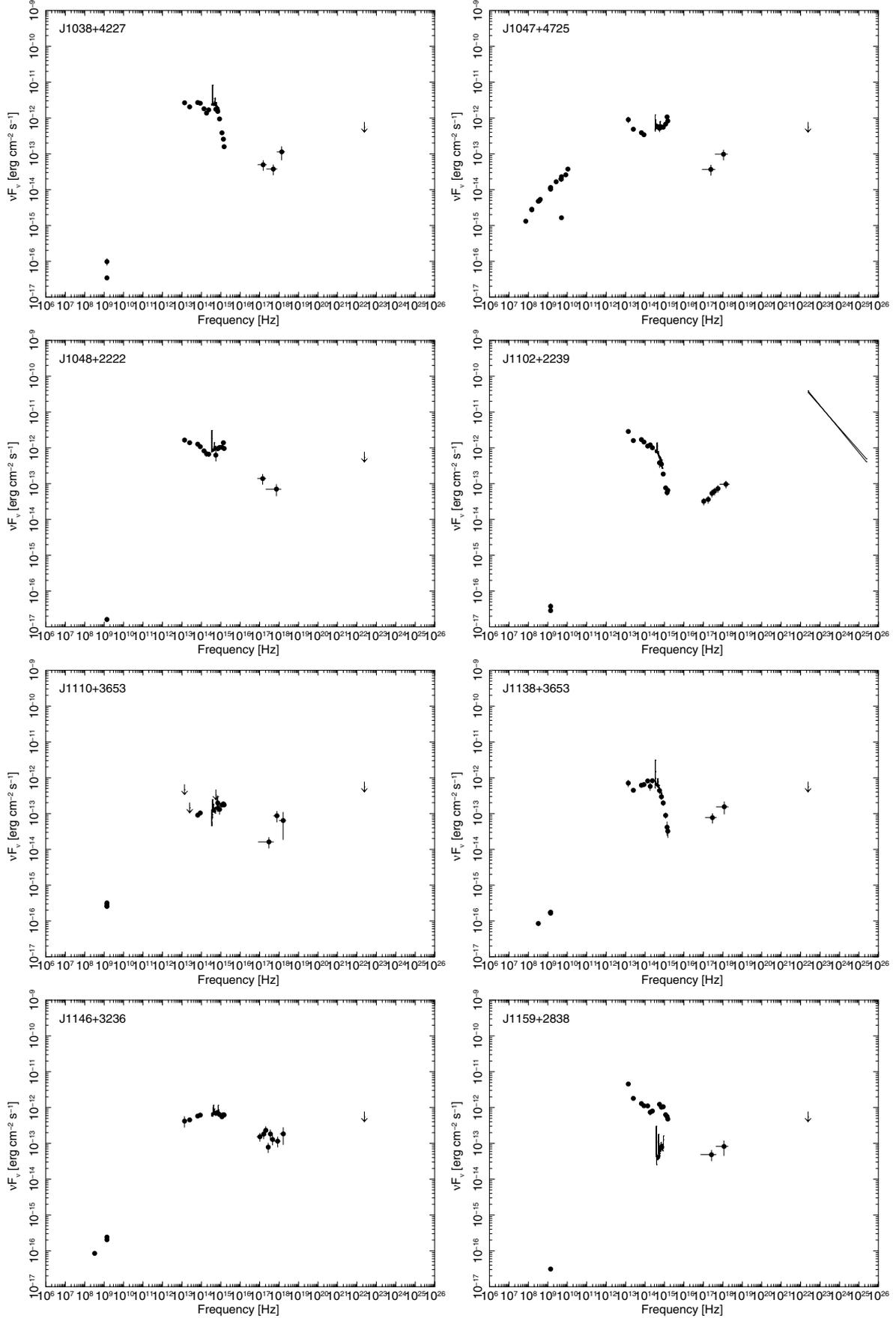


Fig. 10. Spectral energy distributions of the sources in the present sample. Data are corrected for the Galactic absorption. Points refer to detections; arrows are upper limits; the continuous lines are the optical spectra.

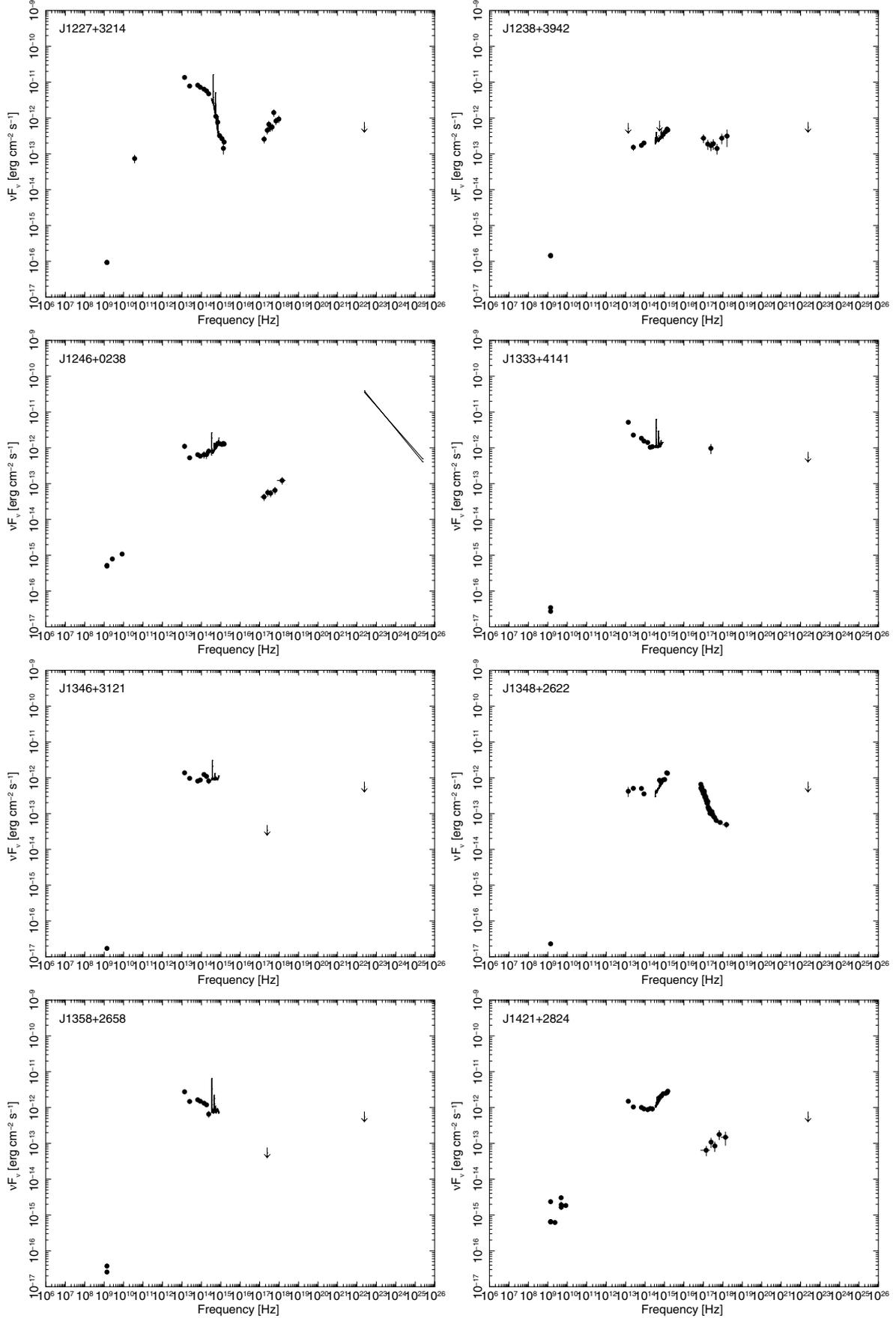


Fig. 11. Spectral energy distributions of the sources in the present sample. Data are corrected for the Galactic absorption. Points refer to detections; arrows are upper limits; the continuous lines are the optical spectra.

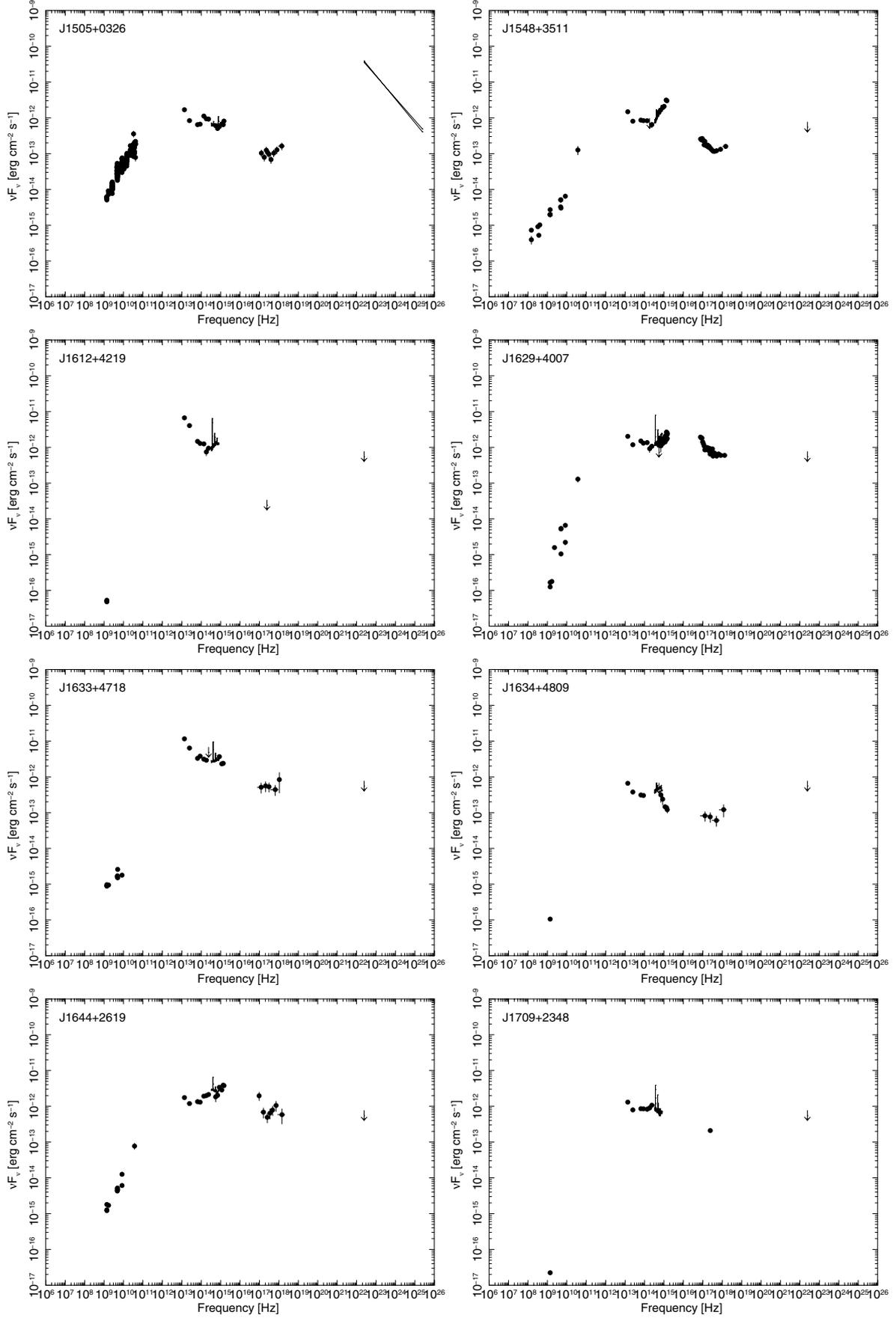


Fig. 12. Spectral energy distributions of the sources in the present sample. Data are corrected for the Galactic absorption. Points refer to detections; arrows are upper limits; the continuous lines are the optical spectra.

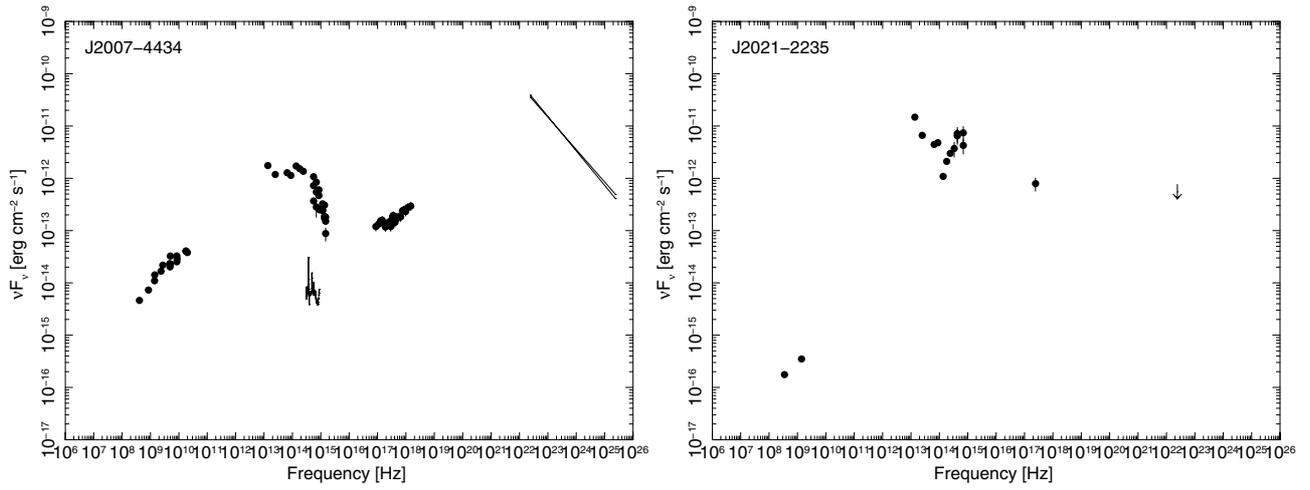


Fig. 13. Spectral energy distributions of the sources in the present sample. Data are corrected for the Galactic absorption. Points refer to detections; arrows are upper limits; the continuous lines are the optical spectra.